A Sightseeing Tour of mm Interferometry

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Towards Higher Resolution:
I. Problem

Telescope resolution:

- $\sim \lambda/D$;
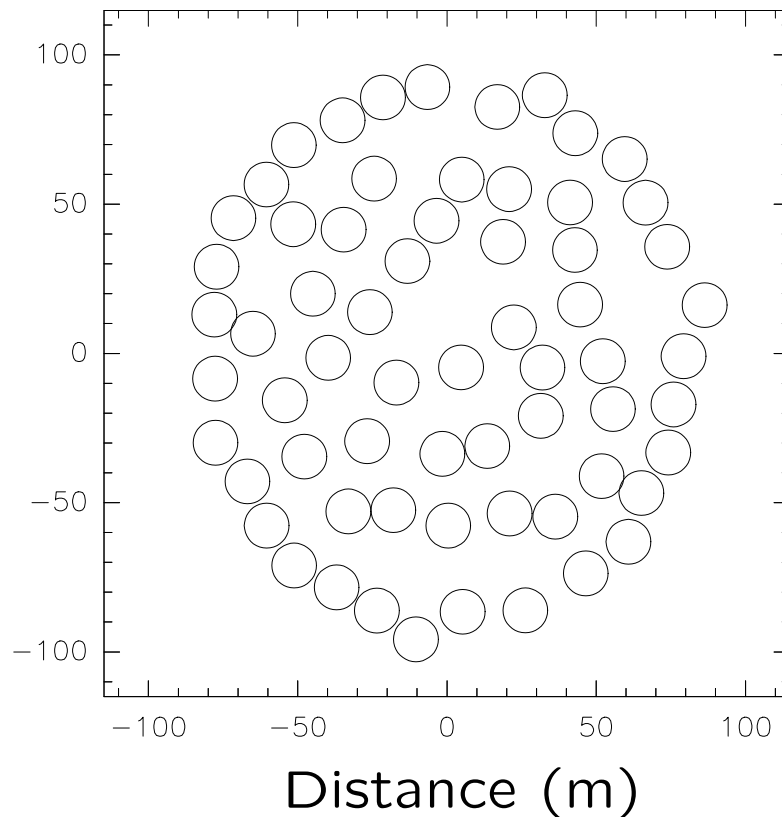- IRAM-30m: $\sim 11''$ @ $1\,\mathrm{mm}$.

Needs to:

- increase D;
- increase precision of telescope positionning;
- keep high surface accuracy.
- $\Rightarrow$ Technically difficult (perhaps impossible?).

Towards Higher Resolution:
II. Solution

Aperture Synthesis: Replacing a single large telescope by a collection of small telescope "filling" the large one.
$\Rightarrow$ Technically difficult but feasible.

ALMA

Distance (m)

Vocabulary and notations:

Baseline Line segment between two antenna.

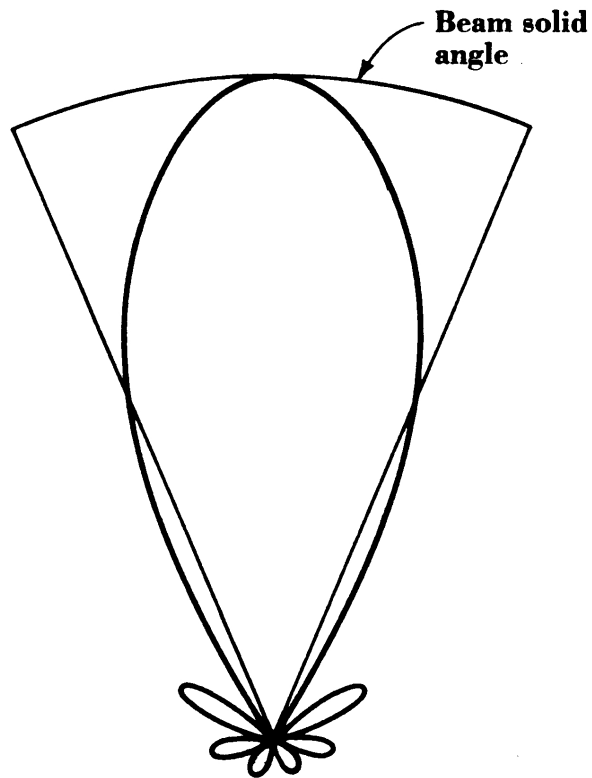b_{ij} Baseline length between antenna i and j.

Configuration Antenna layout (*e.g.* compact configuration).

D configuration size (*e.g.* 150 m).

Primary beam resolution of one antenna (*e.g.* 27″ @ 1 mm).

Synthesized beam resolution of the array (*e.g.* 2″ @ 1 mm).

Parenthesis: PSF = Diffraction Pattern = Beam Pattern

Single-Dish sensitivity
in polar coordinates.

Combination of:

- Antenna properties;
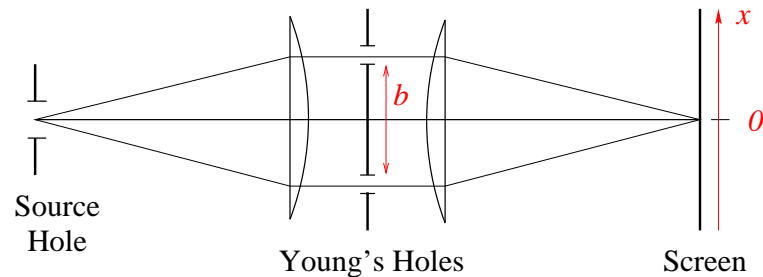- Optical system (*i.e.* how the waves are feeding the receiver).

Typical kind:

Optic/IR Airy function;

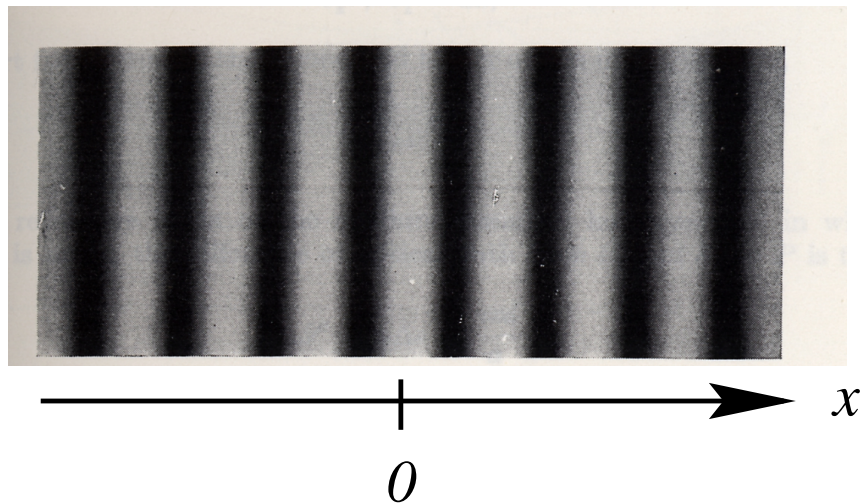**Radio** Gaussian function.

(Lecture by P. Hily-Blant)

Young's Experiment

Setup

Source Hole — Young's Holes — Screen

Lens $\Rightarrow$ Fraunhofer conditions (*i.e.* Plane waves as if the source were placed at infinity).

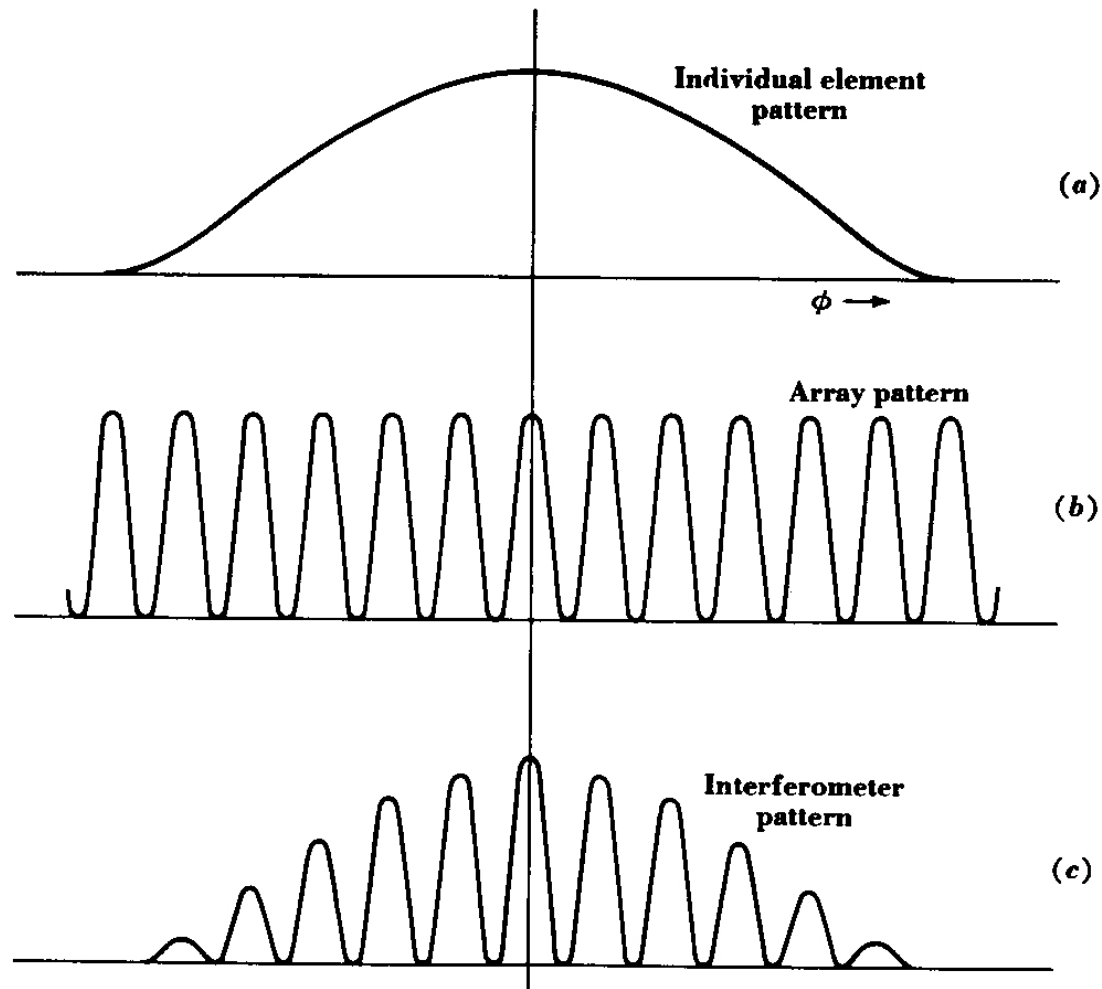## Obtained image of interference: fringes

$$I(x) = I_1 + I_2 + 2\sqrt{I_1 I_2}\cos\left(\frac{bx}{\lambda}\right)$$

with
$$\begin{cases}
\lambda \text{ Source wavelength;} \\
b \text{ Distance between the two Young's holes;} \\
x \text{ Distance from the optical center on the screen.}
\end{cases}$$

Effect of the Antenna Diffraction Pattern

(a) Individual element pattern

$\phi \longrightarrow$

(b) Array pattern

(c) Interferometer pattern

$$I(x) = B(x).\left\{I_1 + I_2 + 2\sqrt{I_1 I_2}\cos\left(\frac{bx}{\lambda}\right)\right\}$$
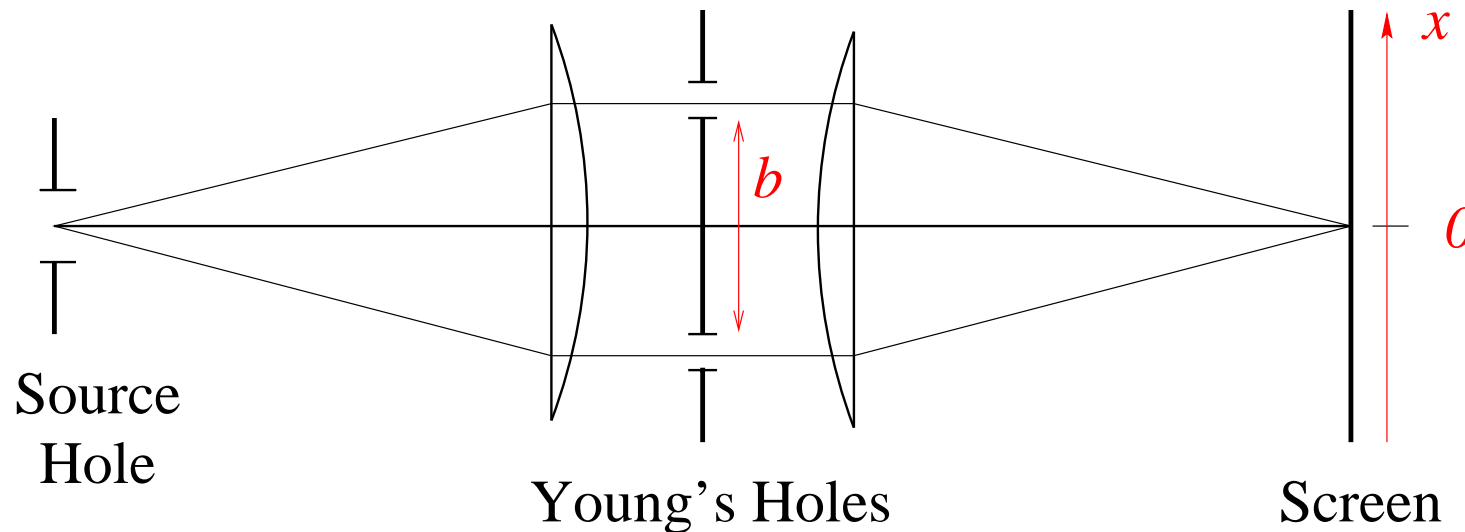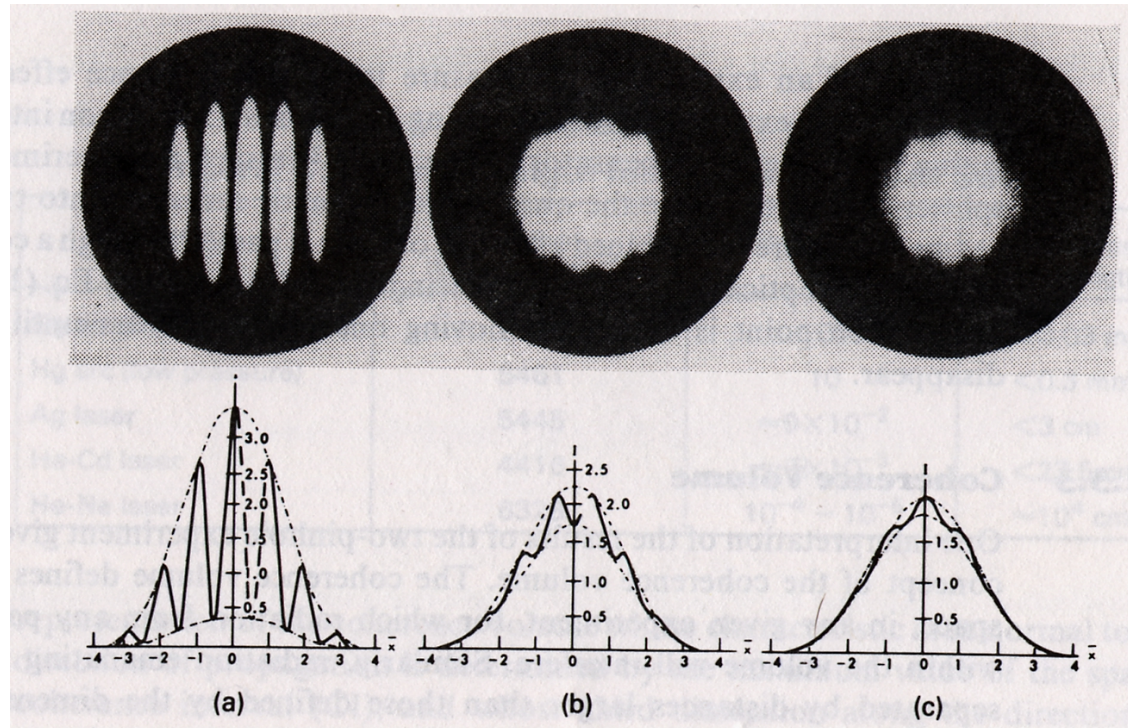
Effect of the Source Hole Size:
I. Description

Hypothesis: Monochromatic source (but not a laser).

Description:

- The Source Hole Size is increased.

- Everything else is kept equal.

Effect of the Source Hole Size:
II. Results

Fringes disappear! $\Rightarrow$ $\begin{cases} \text{Fringe contrast is linked to the} \\ \text{spatial properties of the source.} \end{cases}$

$$I(x) = I_1 + I_2 + 2\sqrt{I_1 I_2}|C|\cos\left(\frac{bx}{\lambda} + \phi_C\right) \text{ with } |C| = \frac{I_{\max} - I_{\min}}{I_{\max} + I_{\min}}$$
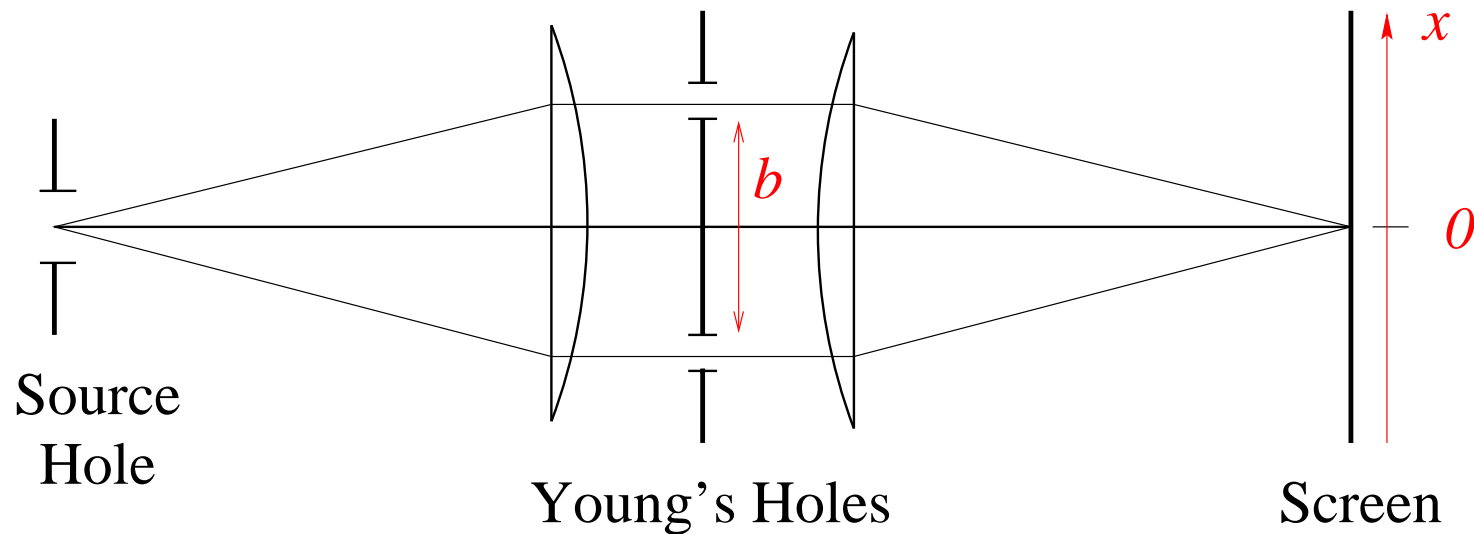
Effect of the Distance Between Young's Holes:
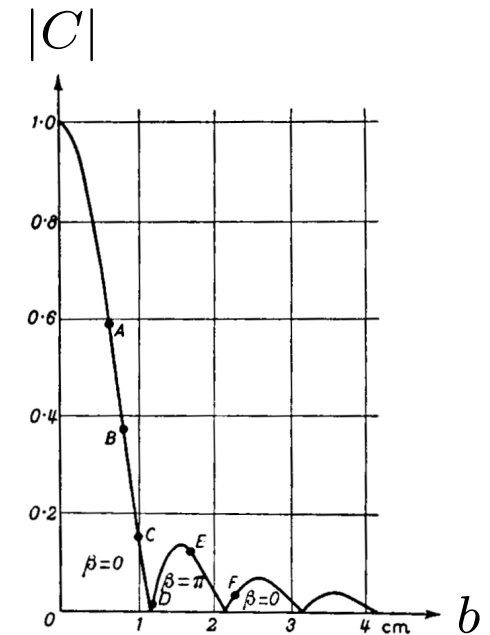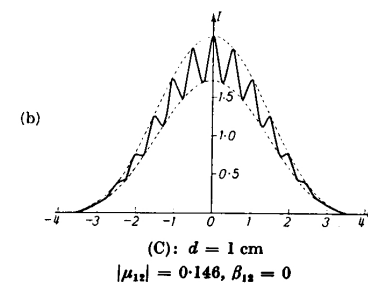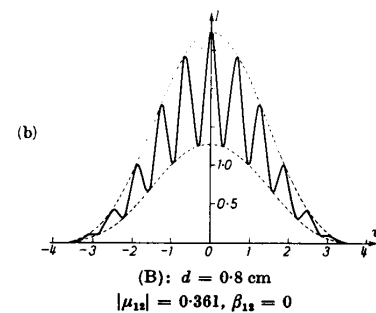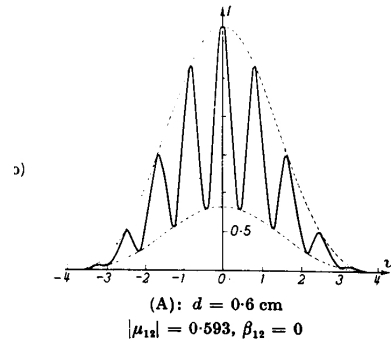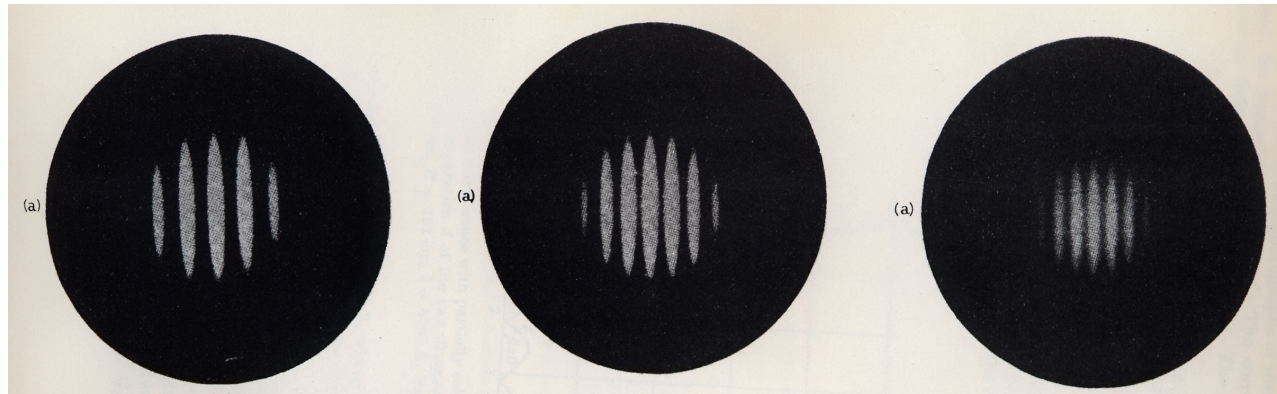## I. Description

Hypothesis:

- Monochromatic source (but not a laser).

- The source hole is a circular disk.

Description:

- The distance between the two Young's holes is increased.

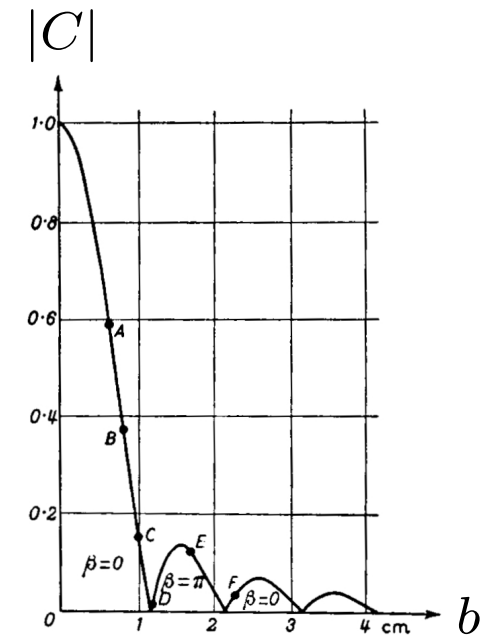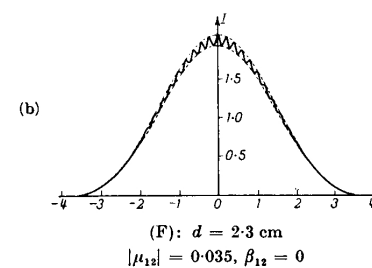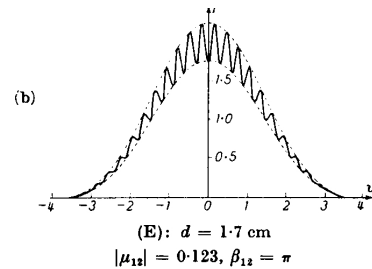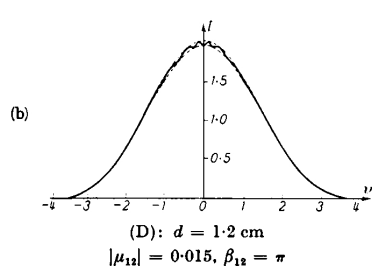- Everything else is kept equal (in particular the hole size).

Source Hole

b

x

0

Young's Holes

Screen

Effect of the Distance Between Young's Holes:
II. Results

$$(A): \quad d = 0{\cdot}6 \text{ cm} \qquad |\mu_{12}| = 0{\cdot}593, \ \beta_{12} = 0$$

$$(B): \quad d = 0{\cdot}8 \text{ cm} \qquad |\mu_{12}| = 0{\cdot}361, \ \beta_{12} = 0$$

$$(C): \quad d = 1 \text{ cm} \qquad |\mu_{12}| = 0{\cdot}146, \ \beta_{12} = 0$$

$$I(x) = I_1 + I_2 + 2\sqrt{I_1 I_2}\,|C|\cos\!\left(\frac{bx}{\lambda} + \phi_C\right) \text{ with } |C| = \frac{I_{\max} - I_{\min}}{I_{\max} + I_{\min}}$$

Effect of the Distance Between Young's Holes:
II. Results (Continued)

(D): $d = 1 \cdot 2$ cm
$|\mu_{12}| = 0 \cdot 015, \; \beta_{12} = \pi$

(E): $d = 1 \cdot 7$ cm
$|\mu_{12}| = 0 \cdot 123, \; \beta_{12} = \pi$

(F): $d = 2 \cdot 3$ cm
$|\mu_{12}| = 0 \cdot 035, \; \beta_{12} = 0$
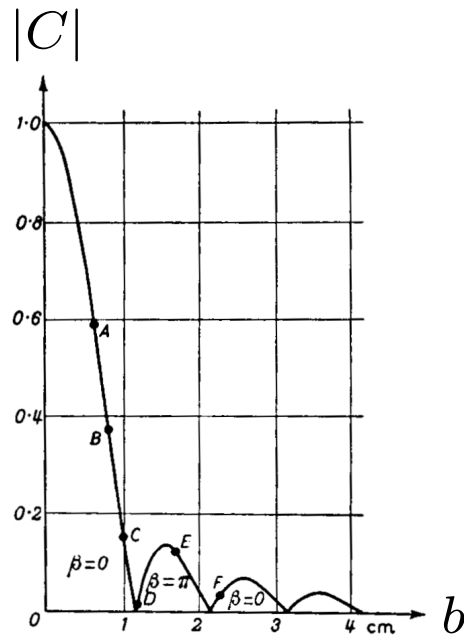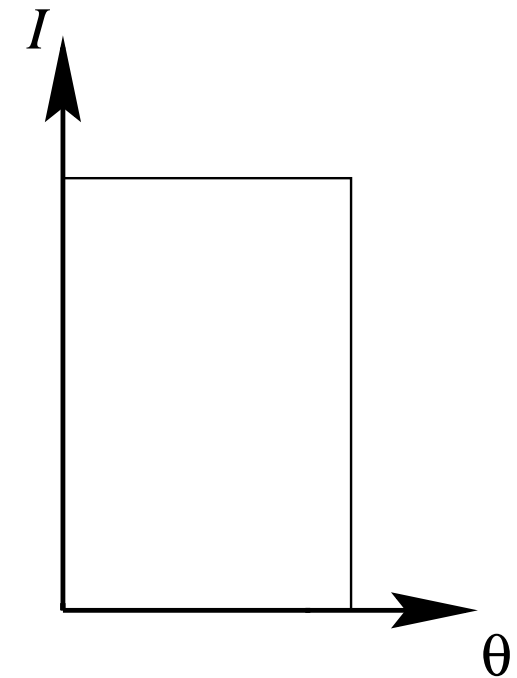
$$I(x) = I_1 + I_2 + 2\sqrt{I_1 I_2}\,|C|\cos\left(\frac{bx}{\lambda} + \phi_C\right) \text{ with } |C| = \frac{I_\text{max} - I_\text{min}}{I_\text{max} + I_\text{min}}$$
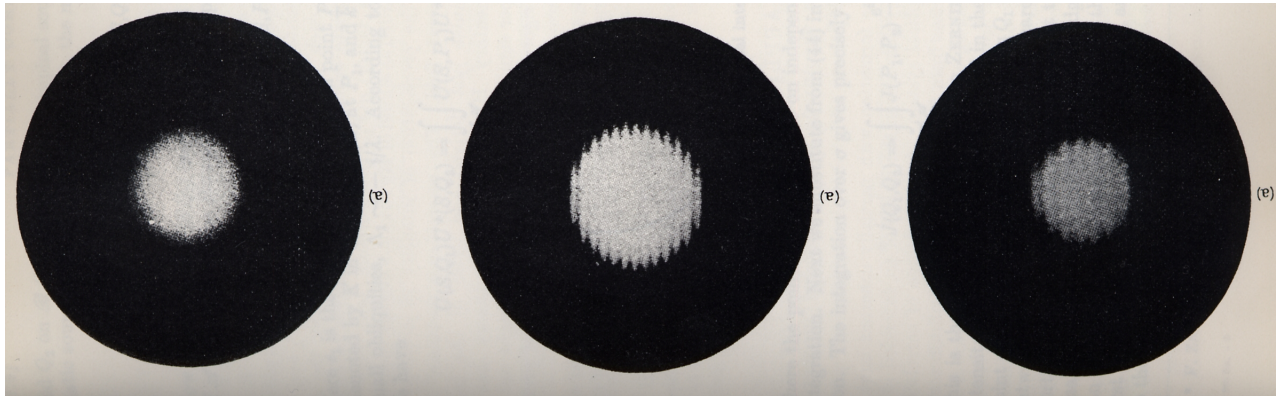
Measured Curve = 2D Fourier Transform of the Source

$$\frac{J_1(b)}{b} \underset{\rightleftharpoons}{\text{2D FT}} \text{Heaviside}(\theta)$$

Source = Uniformly illuminated disk.

Theoretical Basis of the Aperture Synthesis

The van Citter-Zernike theorem

$$V_{ij}(b_{ij}) = C_{ij}(b_{ij}).I_\text{tot} \overset{\text{2D FT}}{\rightleftharpoons} B_\text{primary}.I_\text{source}$$

- Young's holes = Telescopes;

- Signal received by telescopes are combined by pairs;

- Fringe visibilities are measured.

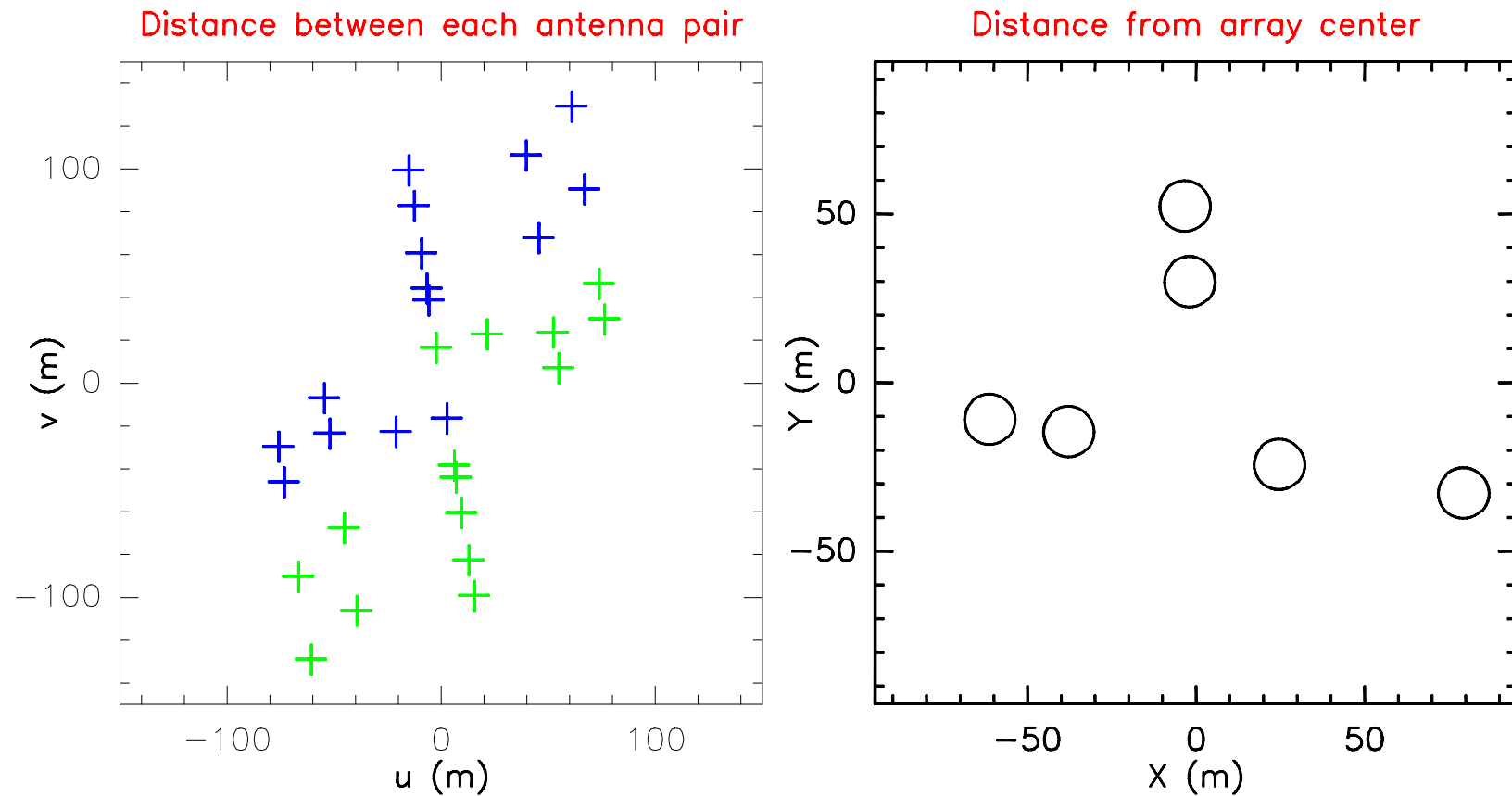$\Rightarrow$ One Fourier component of the source (*i.e.* one visibility) is measured by baseline (or antenna pair).

$\quad\Rightarrow$ Convention: Spatial frequencies are measured in meter.
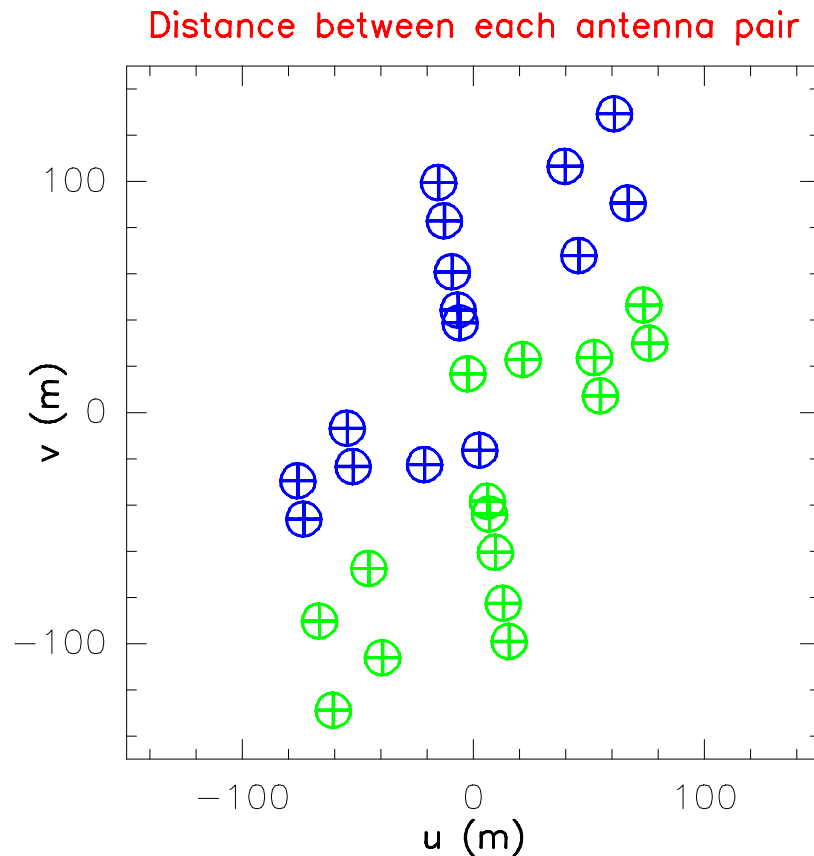
$\quad\quad\Rightarrow$ Each baseline lenght b_{ij} = a spatial frequency.

An Example: The PdBI

Number of baselines: $N(N-1) = 30$ for $N = 6$ antennas.

Convention: Fourier plane $= uv$ plane.

Distance between each antenna pair

Distance from array center

Each Visibility is a Weighted Sum
of the Fourier Components of the Source

Distance between each antenna pair

v (m)

u (m)

$$V_{ij}(b_{ij}) \overset{\text{2D FT}}{\rightleftharpoons} B_{\text{primary}} \cdot I_{\text{source}}$$

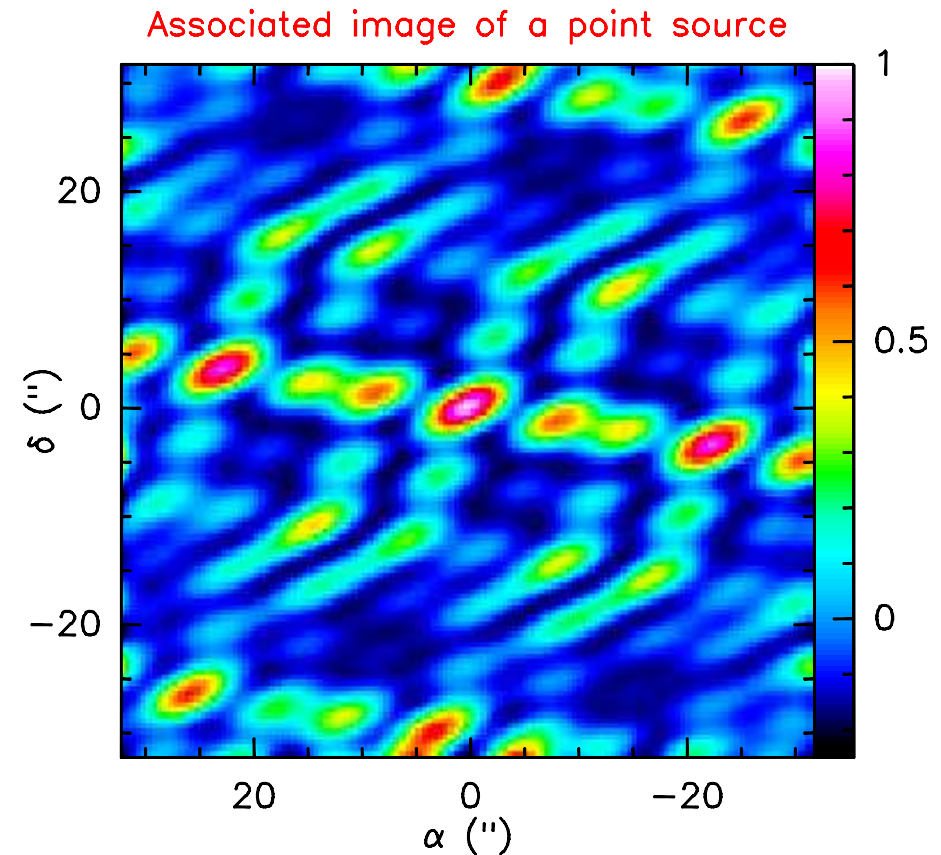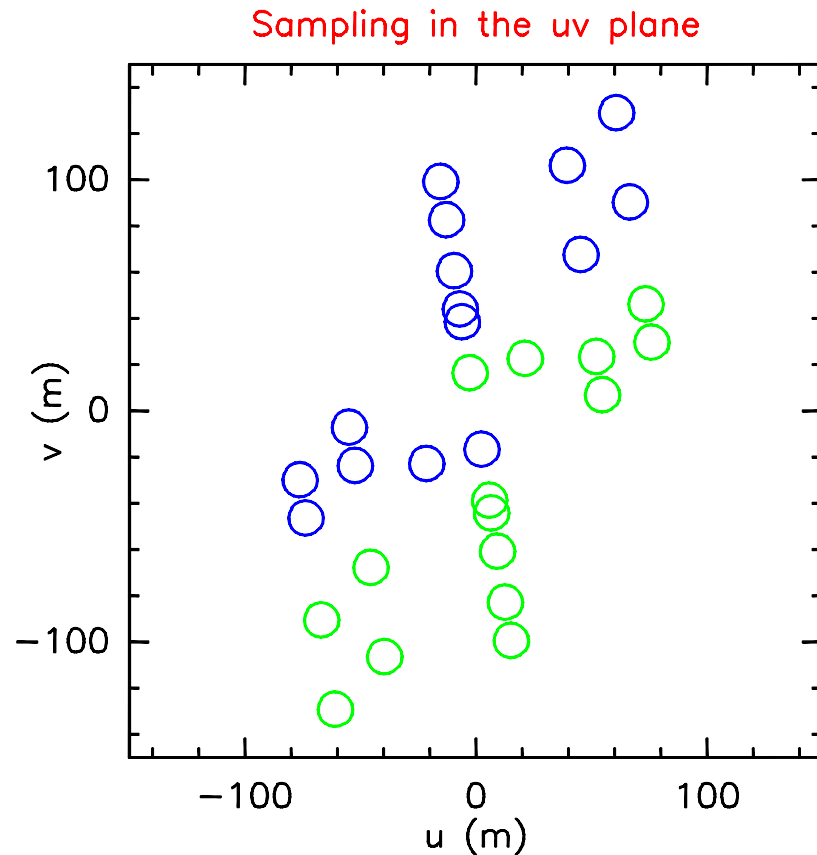$$\text{i.e. } V_{ij}(b_{ij}) = \left\{ \tilde{B}_{\text{primary}} * \tilde{I}_{\text{source}} \right\}(b_{ij})$$

with $\tilde{B}_{\text{primary}}$ a Gaussian of FWHM=15 m.

$$\Rightarrow \begin{cases} \text{Indirect information on the source} \\ \text{(important for mosaicing).} \end{cases}$$

An Example: The PdBI (Cont'd)

Number of baselines: $N(N-1) = 30$ for $N = 6$ antennas.

Convention: Fourier plane $= uv$ plane.

Sampling in the uv plane

Associated image of a point source
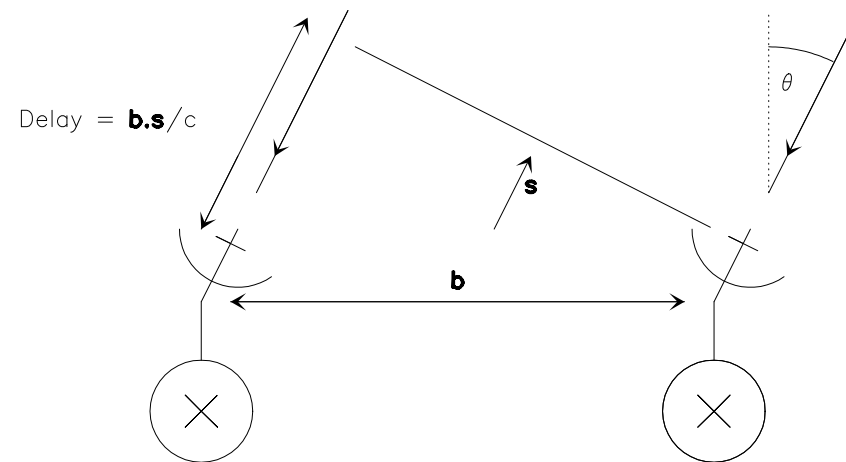
Incomplete uv plane coverage $\Rightarrow$ difficult to make a reliable image
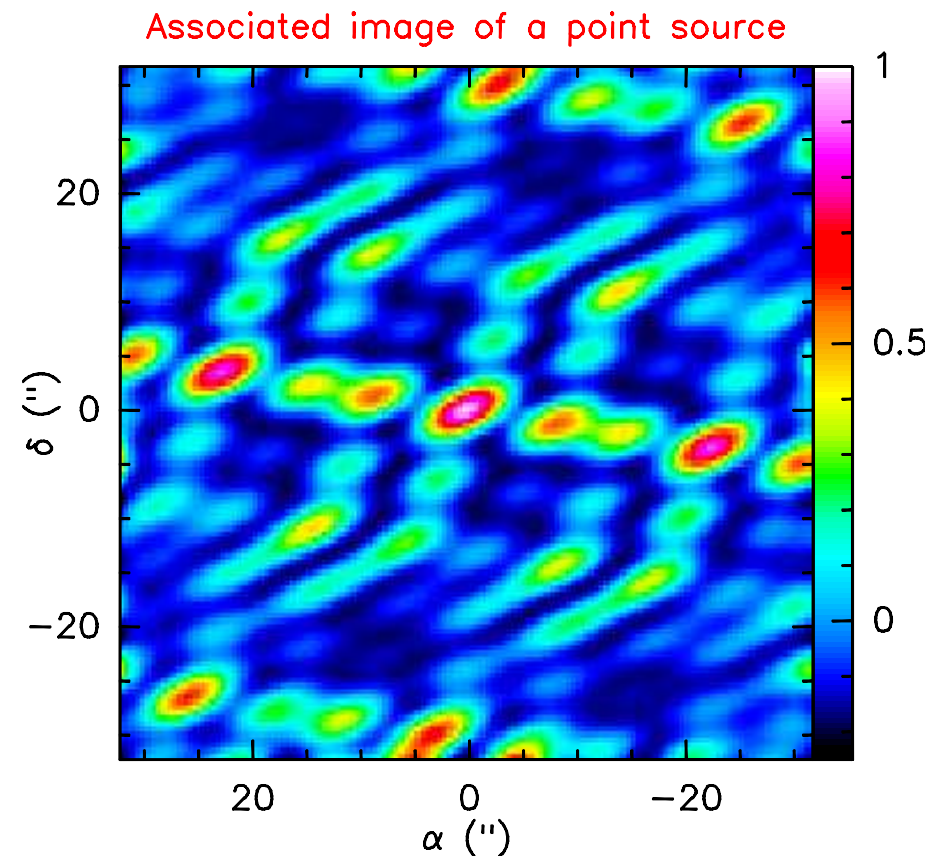(Lectures by A. Castro–Carrizo, J. Pety and F. Gueth).

Precision: Spatial frequencies = baseline lengths projected in a plane perpendicular to the source mean direction.

Delay = $\mathbf{b}.\mathbf{s}/c$
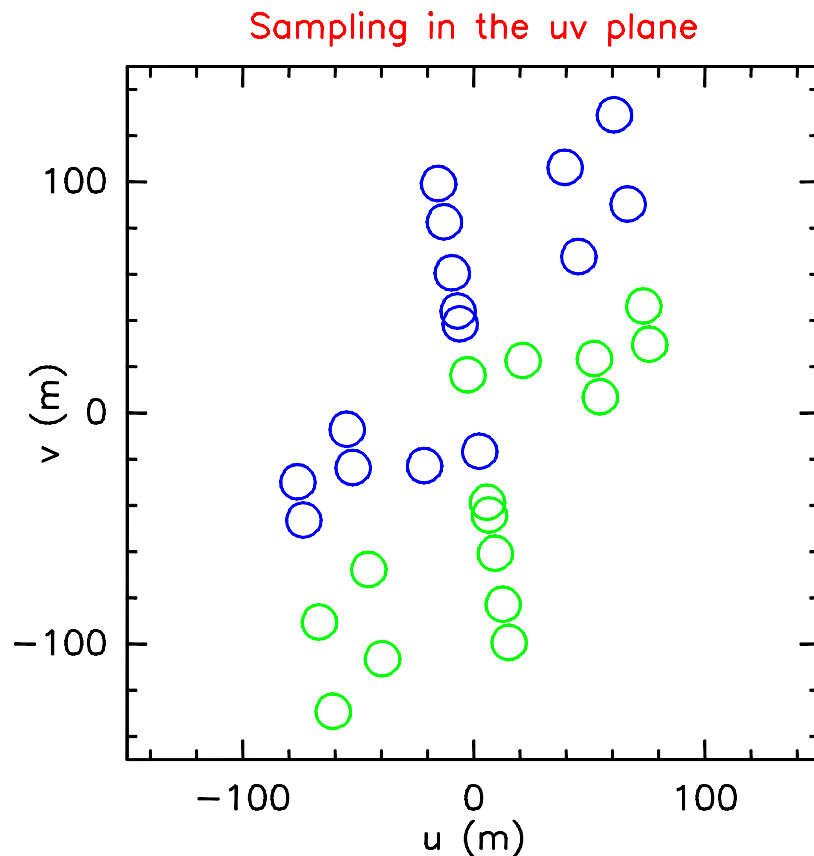
θ

$\mathbf{s}$

$\mathbf{b}$

Earth Rotation and Super Synthesis

Precision: Spatial frequencies = baseline lengths projected in a plane perpendicular to the source mean direction.

Advantage: Possibility to measure different Fourier components without moving antennas!

Sampling in the uv plane

Associated image of a point source

Earth Rotation and Super Synthesis

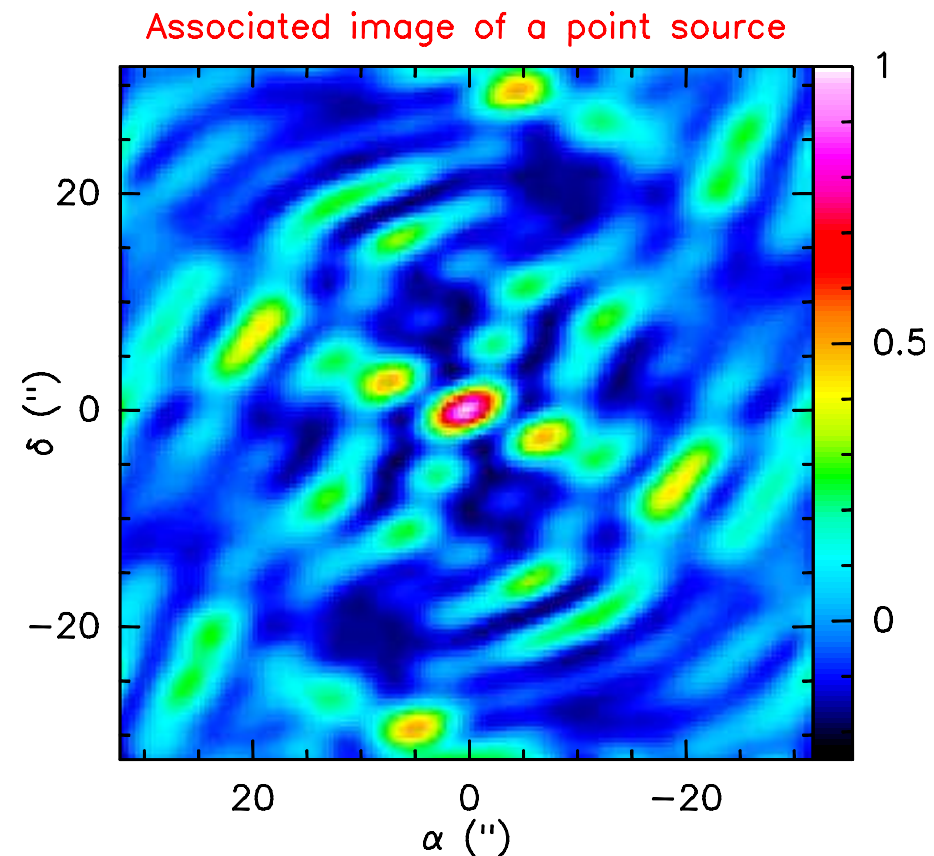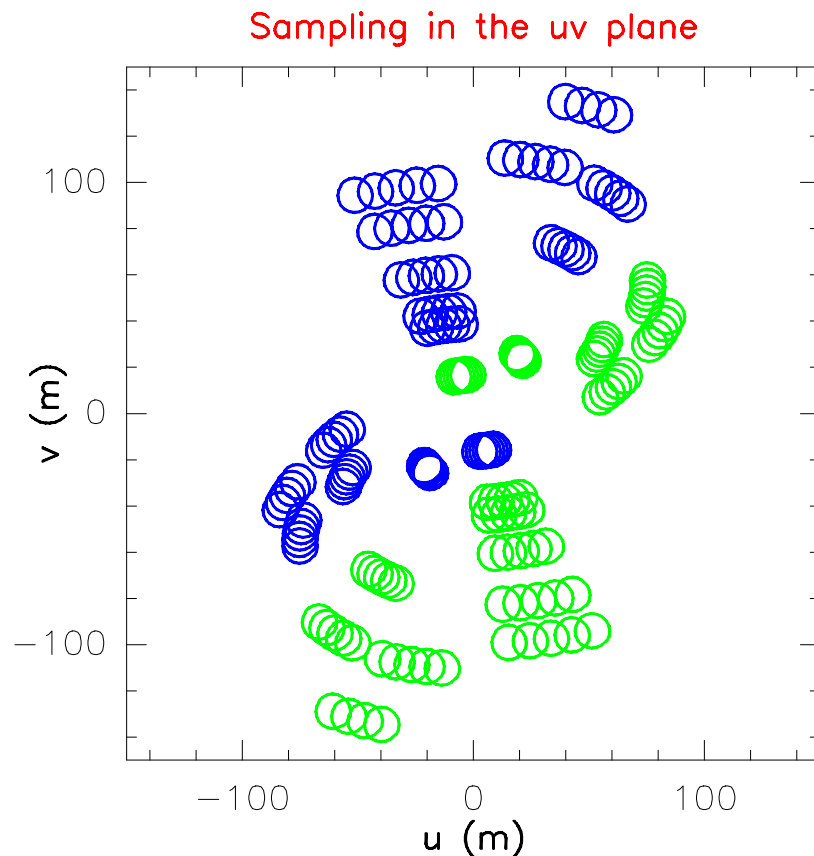Precision: Spatial frequencies = baseline lengths projected in a plane perpendicular to the source mean direction.

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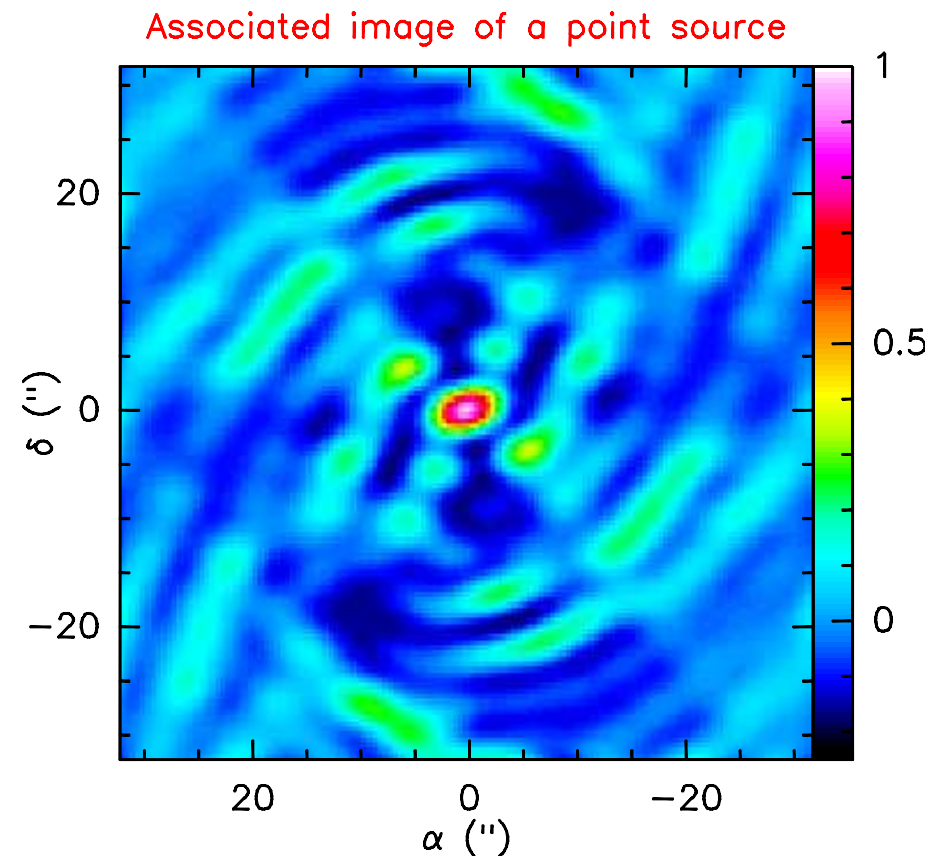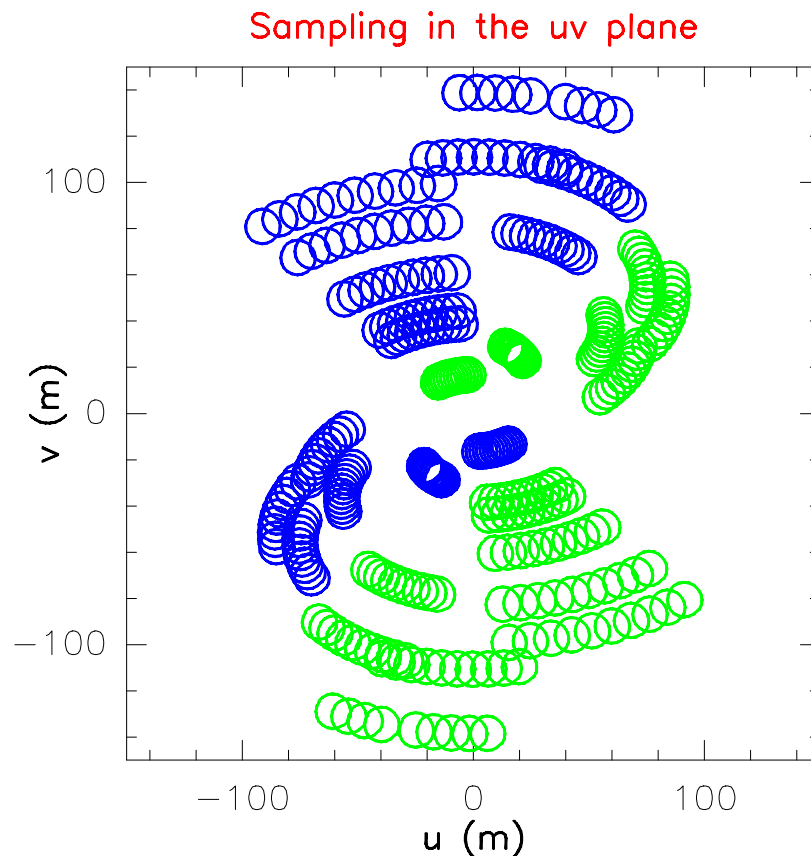Precision: Spatial frequencies = baseline lengths projected in a plane perpendicular to the source mean direction.

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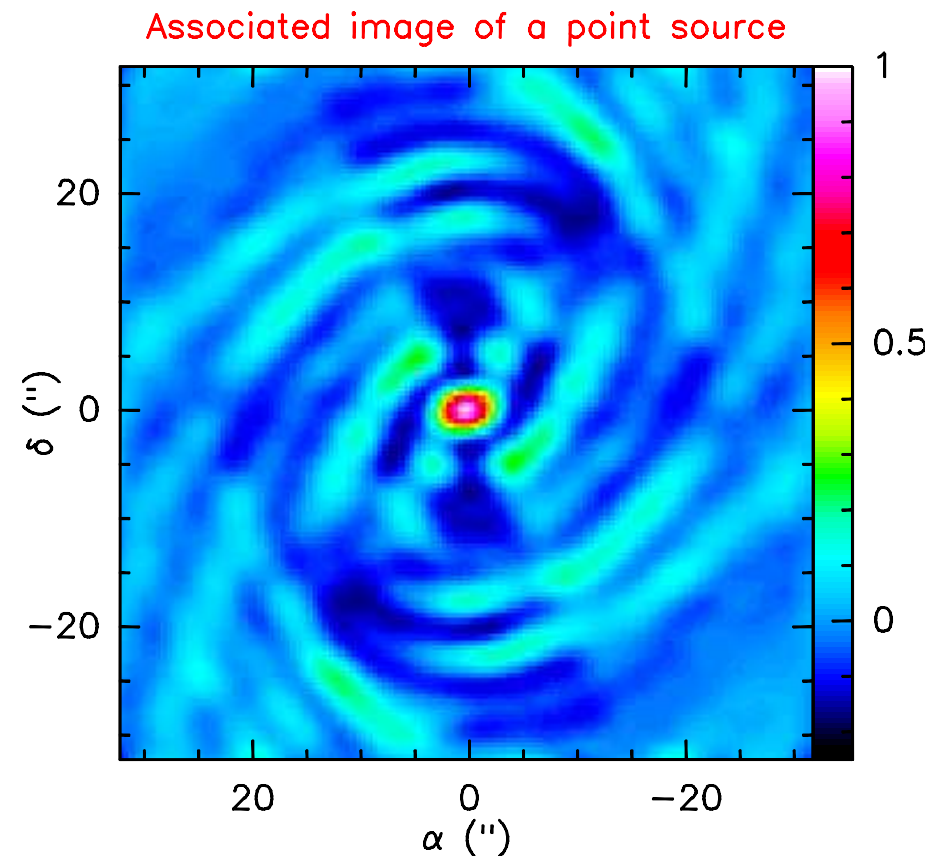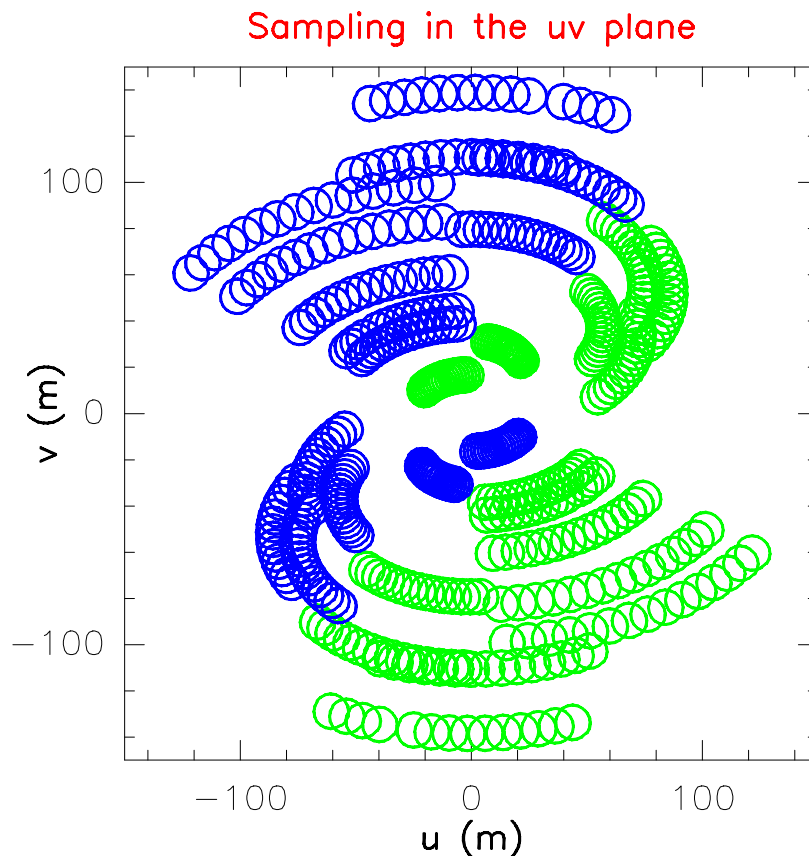Precision: Spatial frequencies = baseline lengths projected in a plane perpendicular to the source mean direction.

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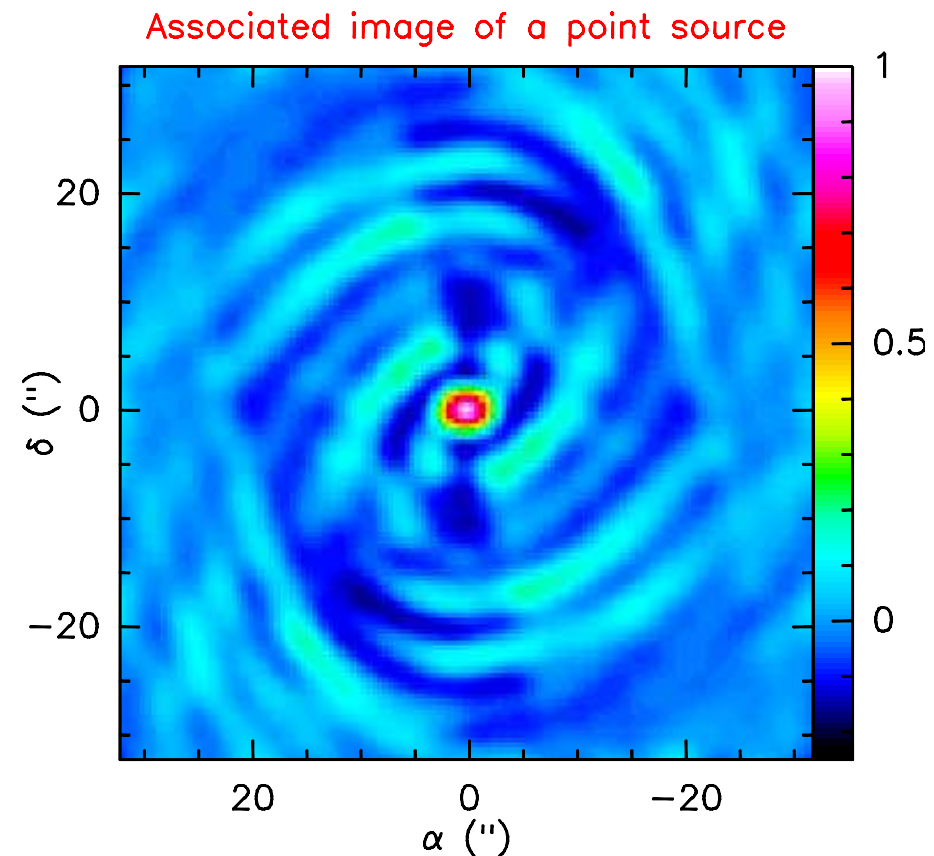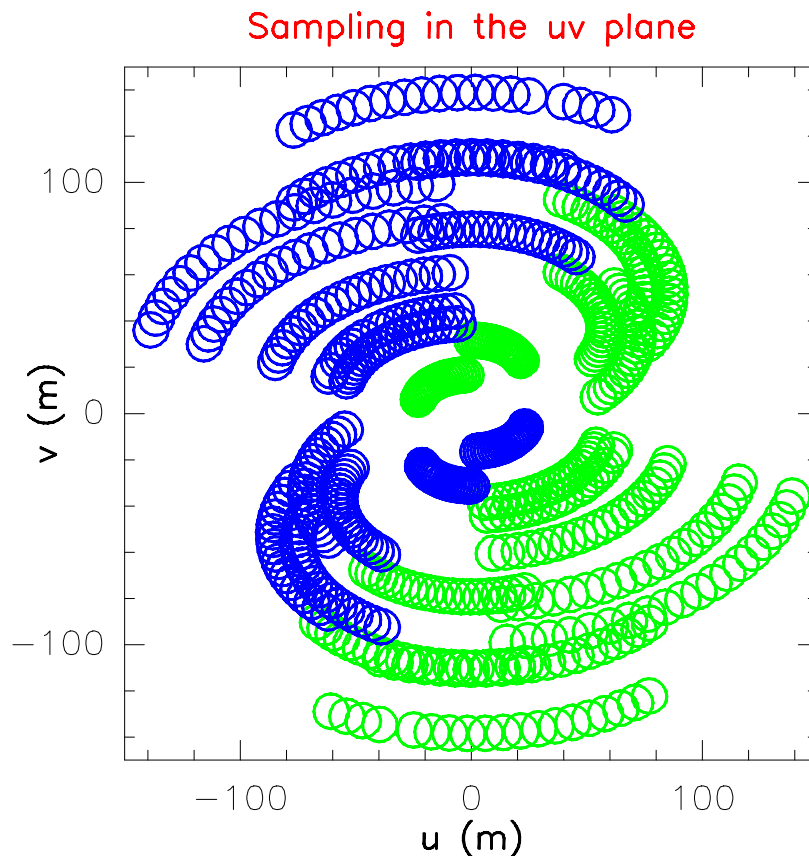Precision: Spatial frequencies = baseline lengths projected in a plane perpendicular to the source mean direction.

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Sampling in the uv plane

Associated image of a point source

Earth Rotation and Super Synthesis

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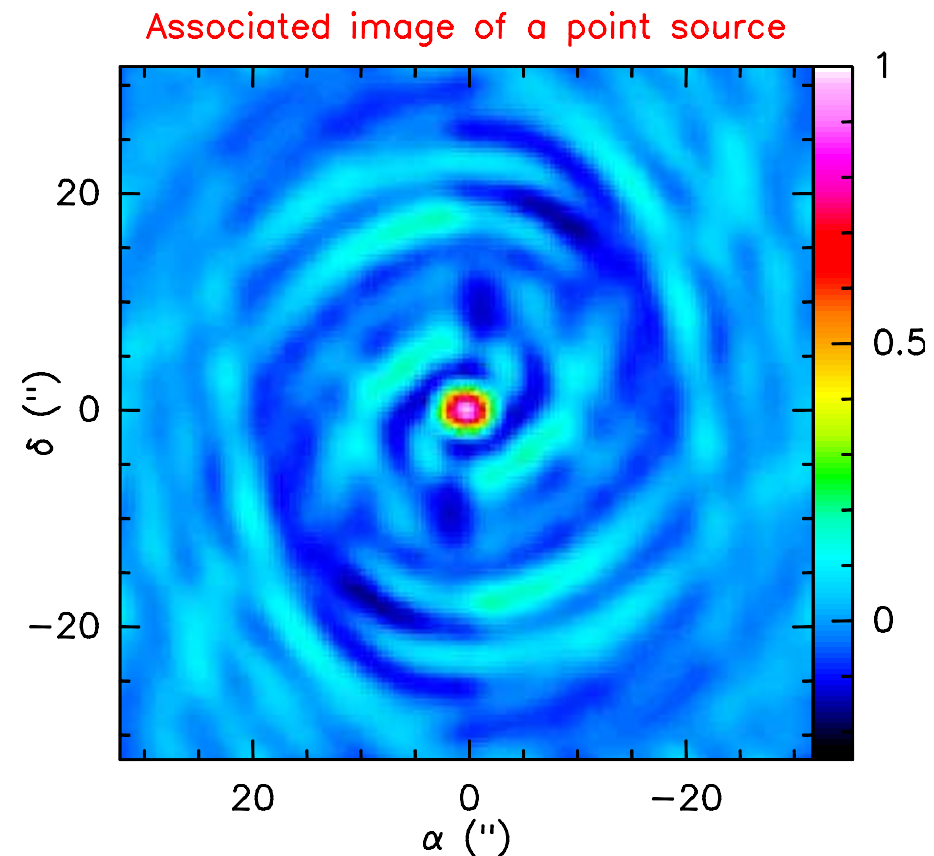Advantage: Possibility to measure different Fourier components without moving antennas!

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Associated image of a point source

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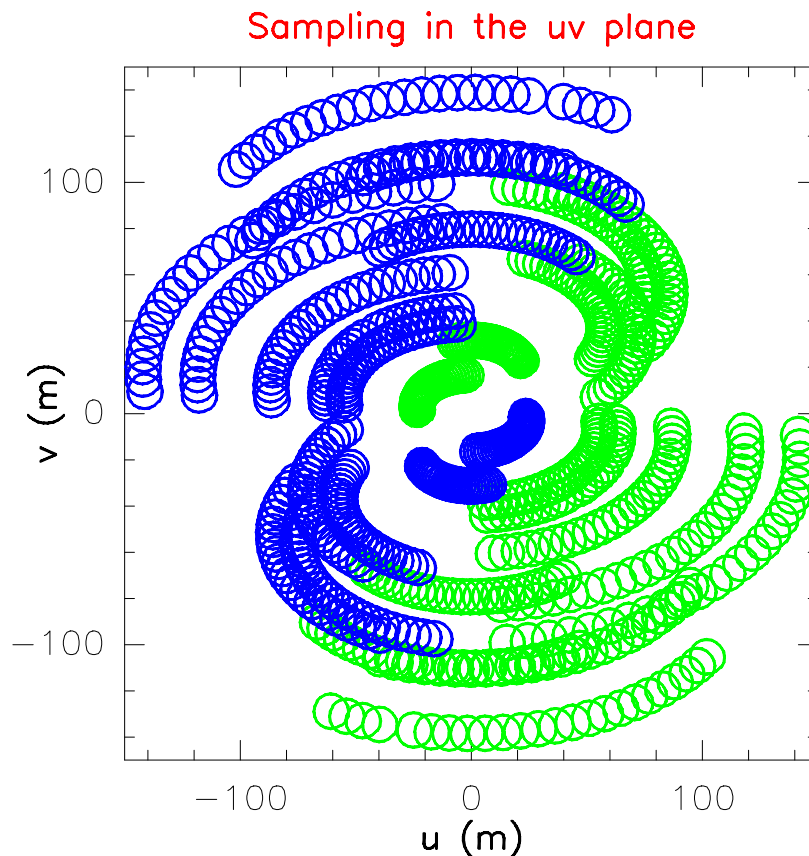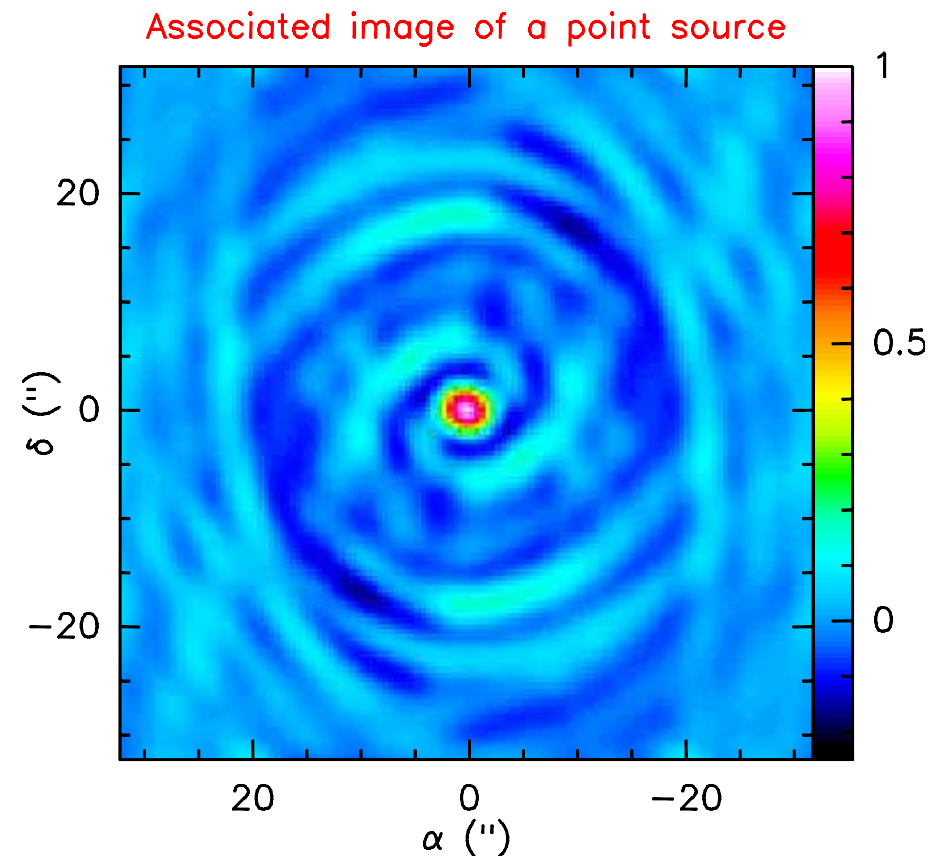Advantage: Possibility to measure different Fourier components without moving antennas!

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Earth Rotation and Super Synthesis

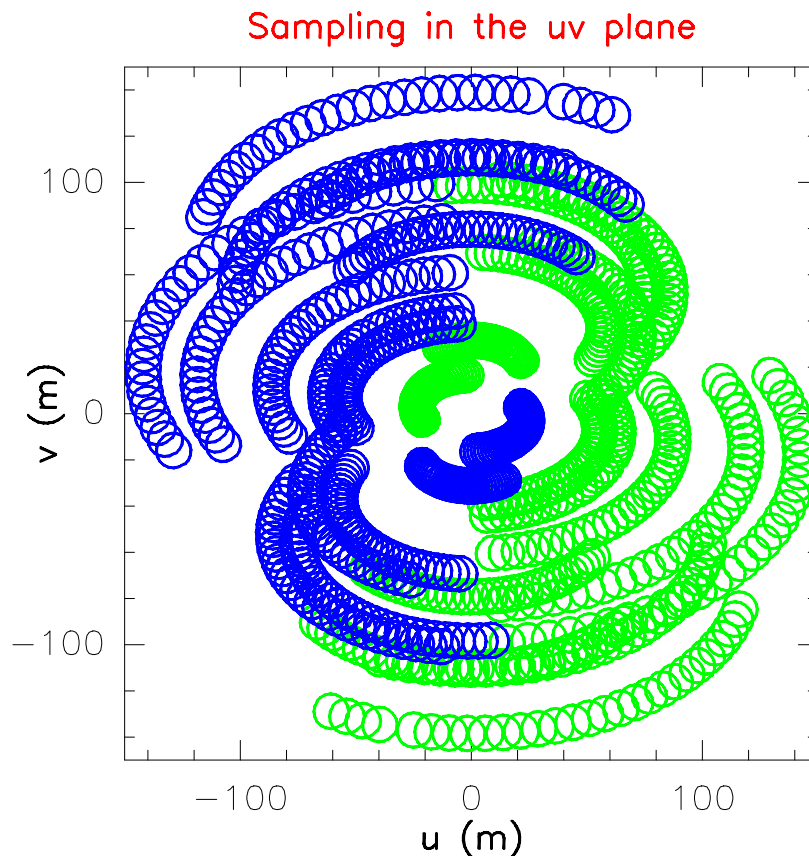Precision: Spatial frequencies = baseline lengths projected in a plane perpendicular to the source mean direction.

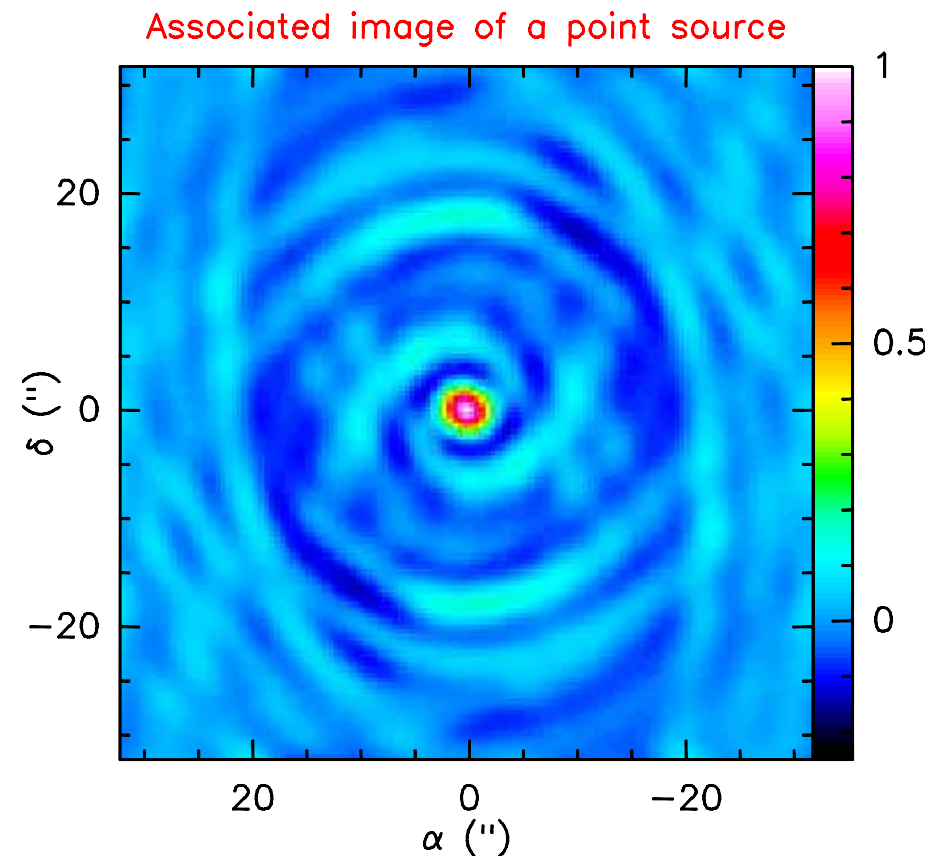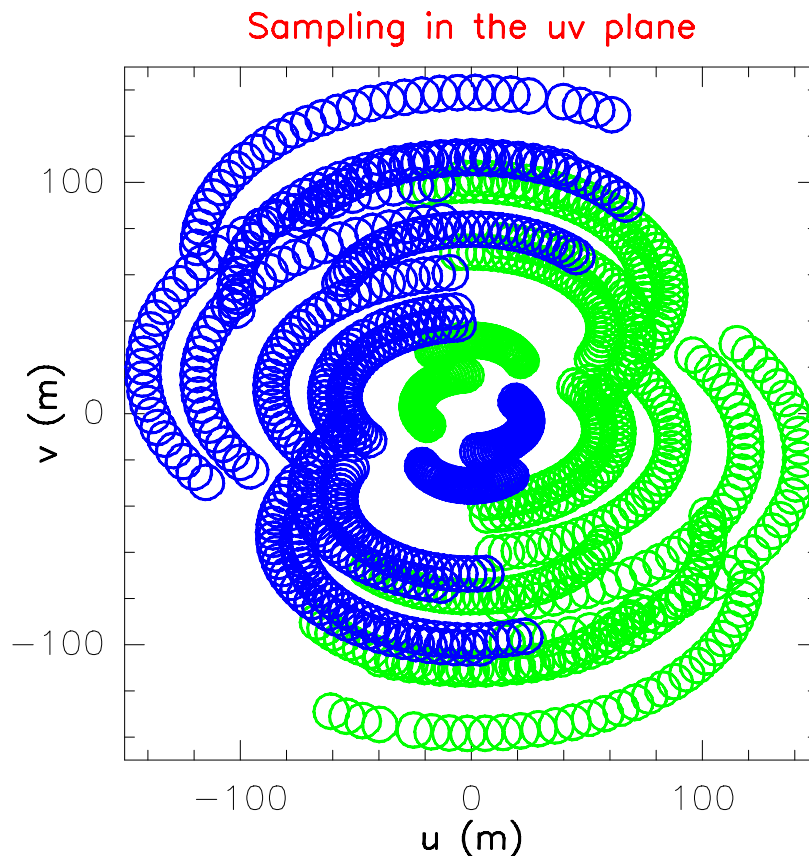Advantage: Possibility to measure different Fourier components without moving antennas!

Sampling in the uv plane

Associated image of a point source

Delay Correction: I. Why?

Real life: Source not at zenith.

$\Rightarrow$ $\begin{cases} \text{Wave plane arrives at different} \\ \qquad \text{moment on each antenna.} \end{cases}$

Delay = $\mathbf{b.s}/c$

Temporal coherence:

- $E(t) = E_0 \cos(\omega t + \psi)$
- Temporally Incoherent Source = random phase changes.
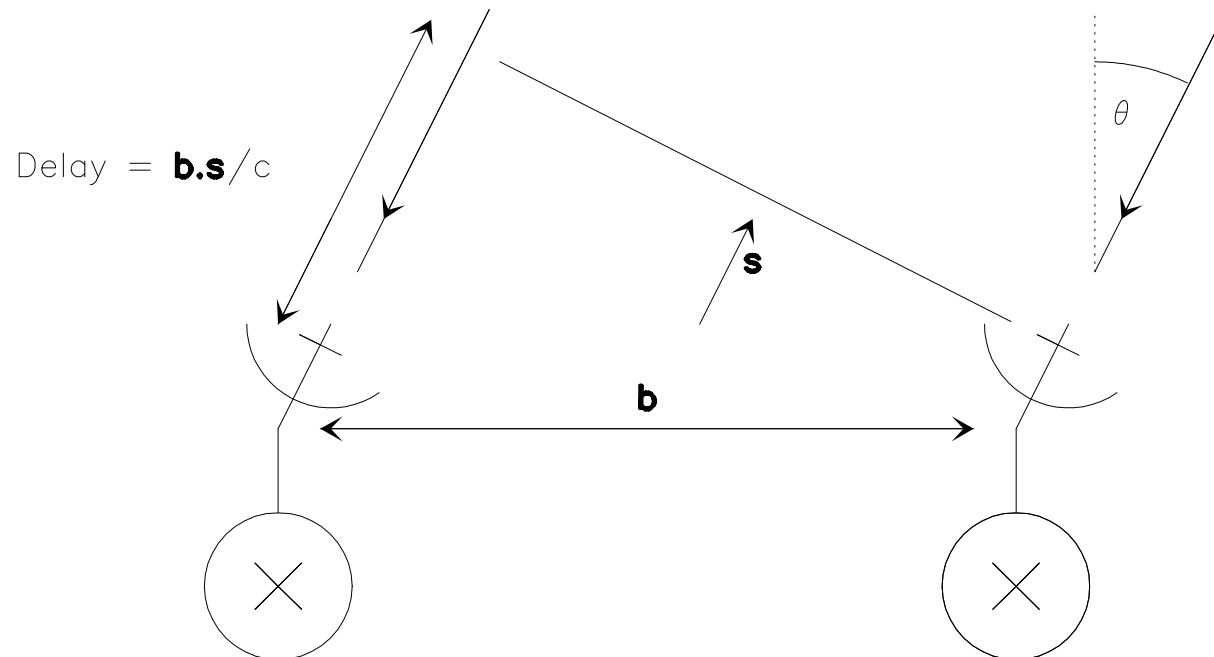- Coherence time: mean time over which wave phase = constant.

$\psi = 0$ $\qquad$ $\psi = 1.5$ $\qquad$ $\psi = 0.5$
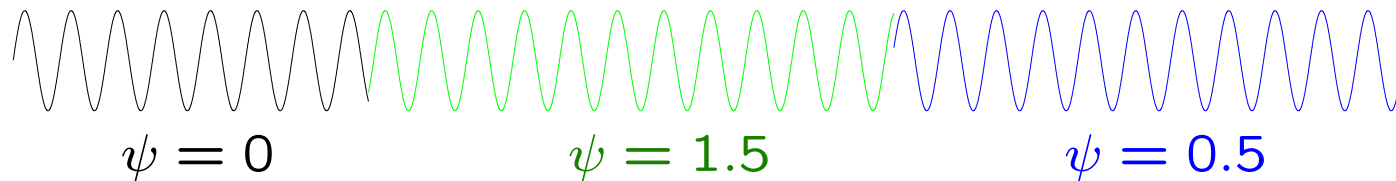
Problem: (Coherence time $\lesssim$ delay) $\Rightarrow$ fringes disappear!
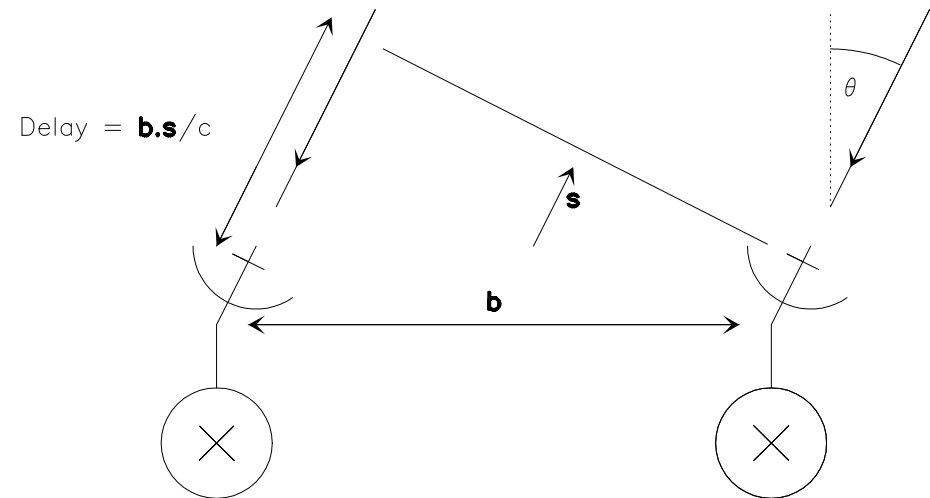
Delay Correction: II. Earth rotation

Earth rotation:

- Advantage: Super synthesis;

- Inconvenient: Delay correction varies with time!

Delay = $\mathbf{b.s}/c$

θ

$\mathbf{s}$

$\mathbf{b}$

Delay Correction: III. Finite Bandwidth

Real life: Observation of finite bandwidth.

$\Rightarrow$ polychromatic light.

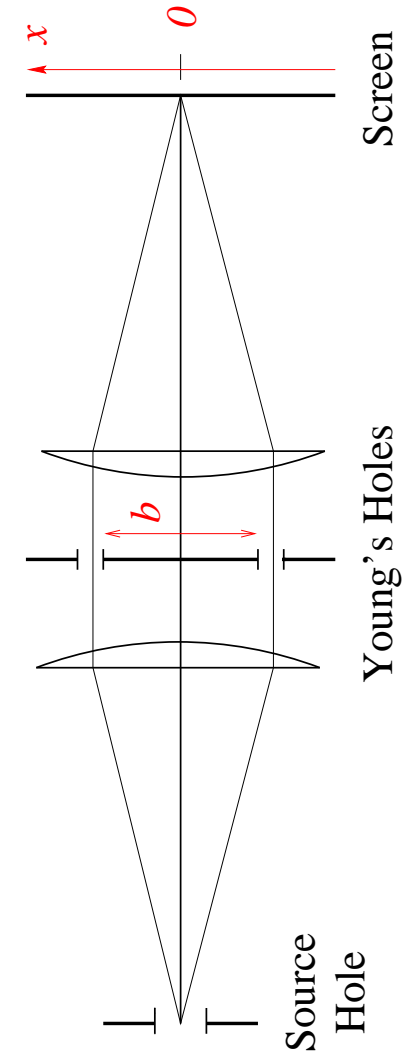Perfect delay correction

$\Rightarrow$ White fringes in 0.

Delay Correction: III. Finite Bandwidth

Real life: Observation of finite bandwidth.
⇒ polychromatic light.

Perfect delay correction
⇒ White fringes in 0.

Worse and worse delay correction.
⇒ Translation of the fringe pattern.
⇒ Fringes seem to disappear.

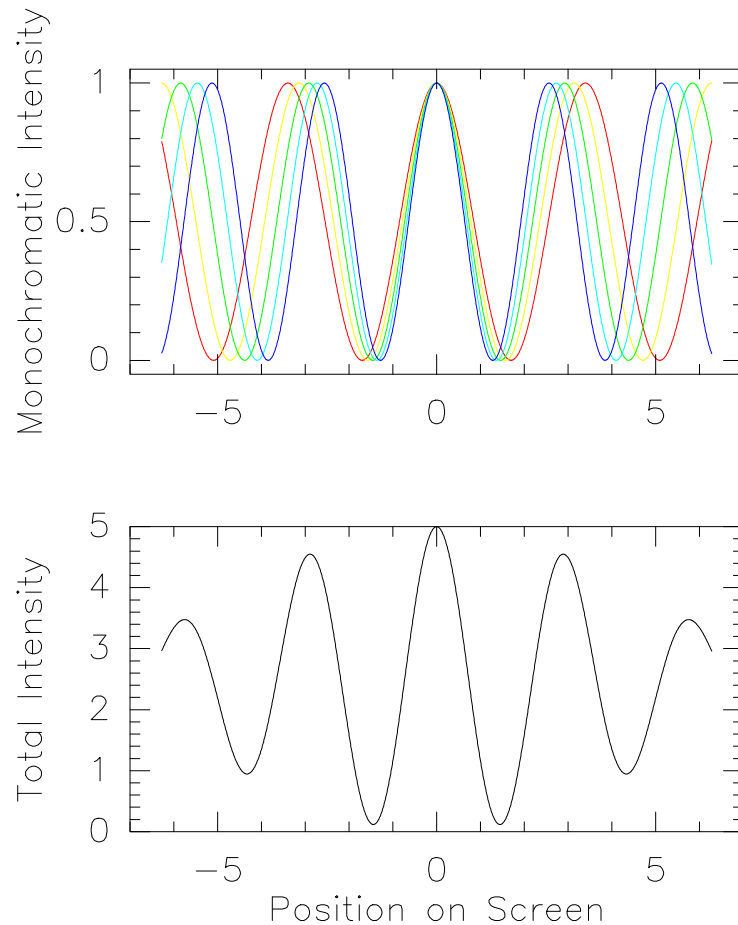# Delay Correction: III. Finite Bandwidth

Real life: Observation of finite bandwidth.

⇒ polychromatic light.
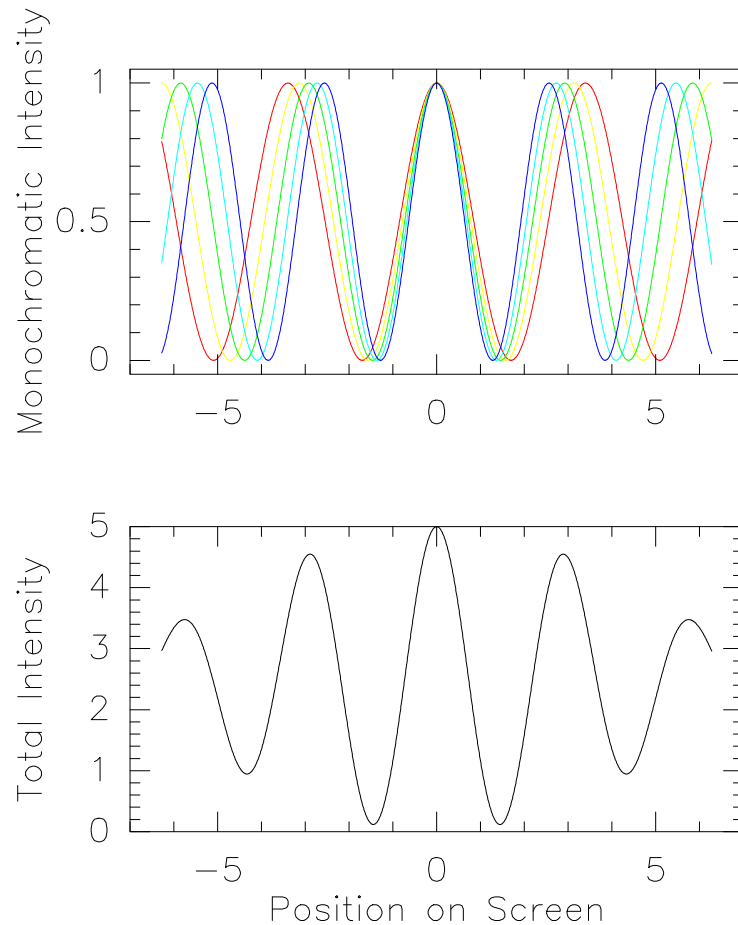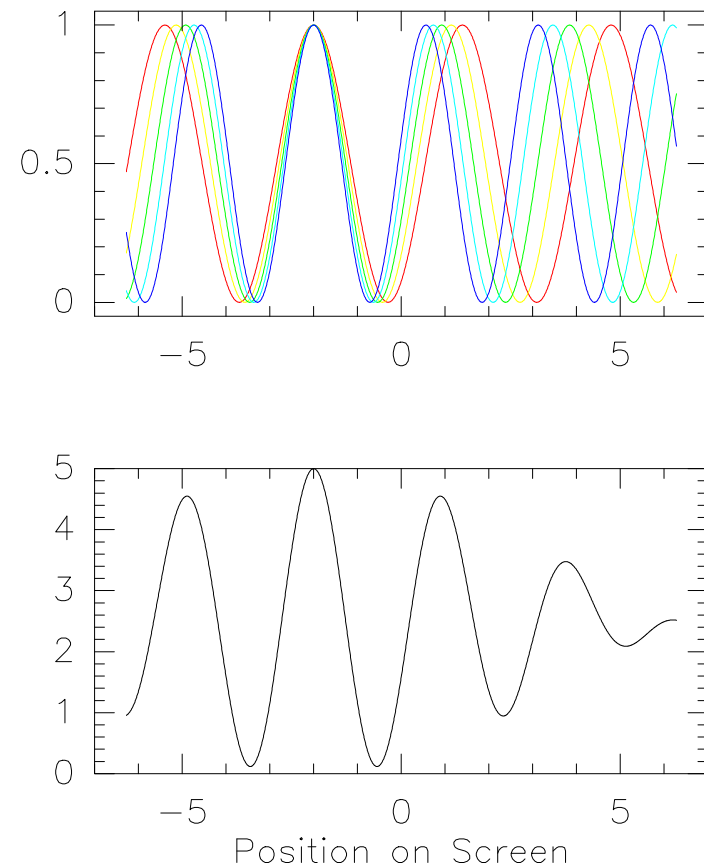
Perfect delay correction
⇒ White fringes in 0.

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Perfect delay correction   Worse and worse delay correction.

⇒ White fringes in 0.   ⇒ Translation of the fringe pattern.

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Worse and worse delay correction.
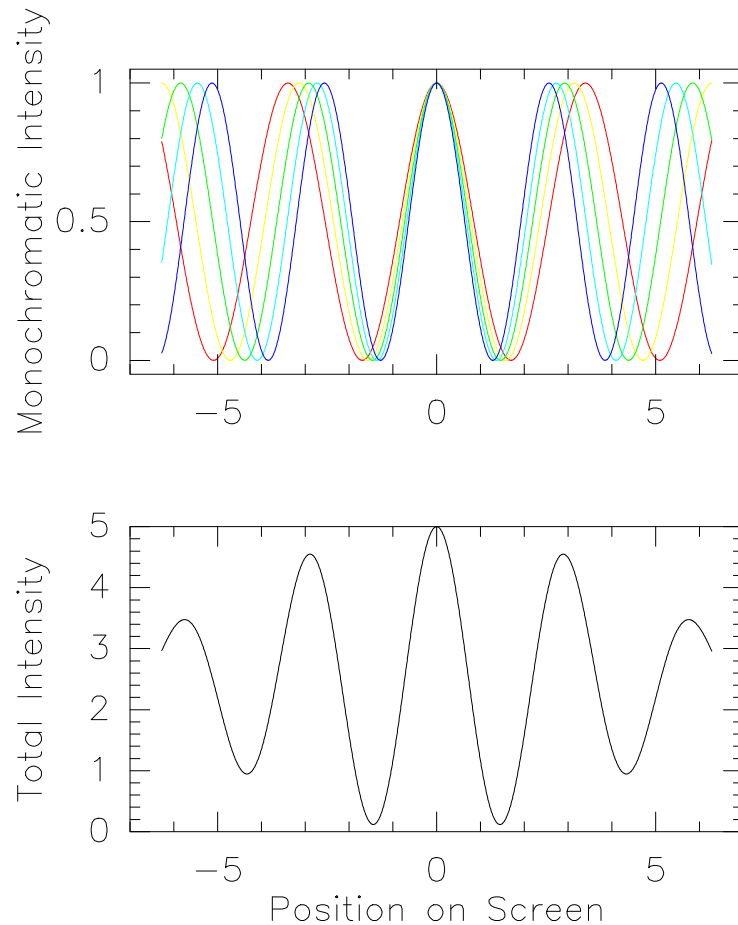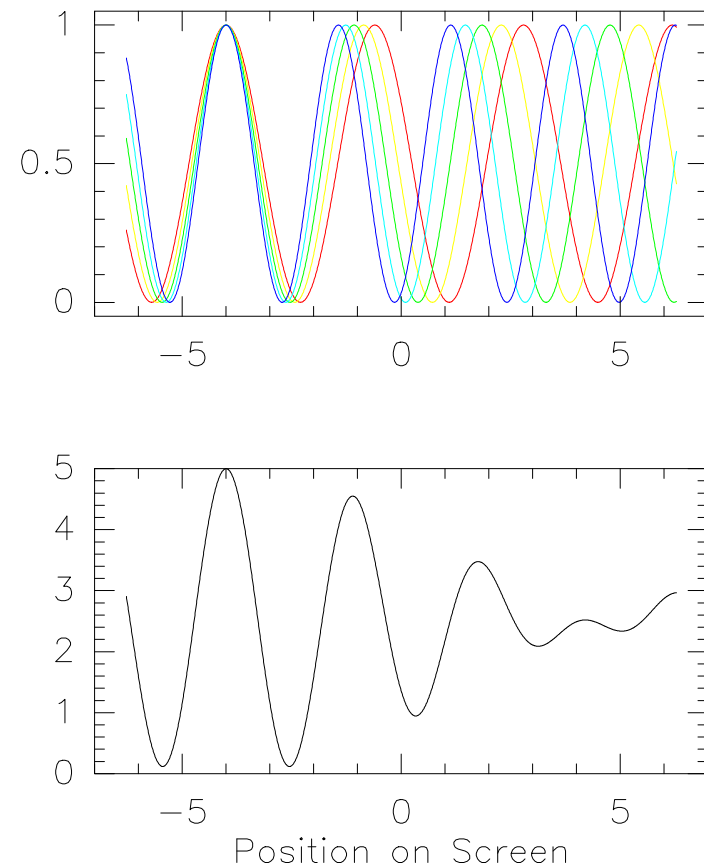⇒ Translation of the fringe pattern.
⇒ Fringes seem to disappear.

Delay Correction: III. Finite Bandwidth

Real life: Observation of finite bandwidth.

⇒ polychromatic light.

Perfect delay correction
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Worse and worse delay correction.
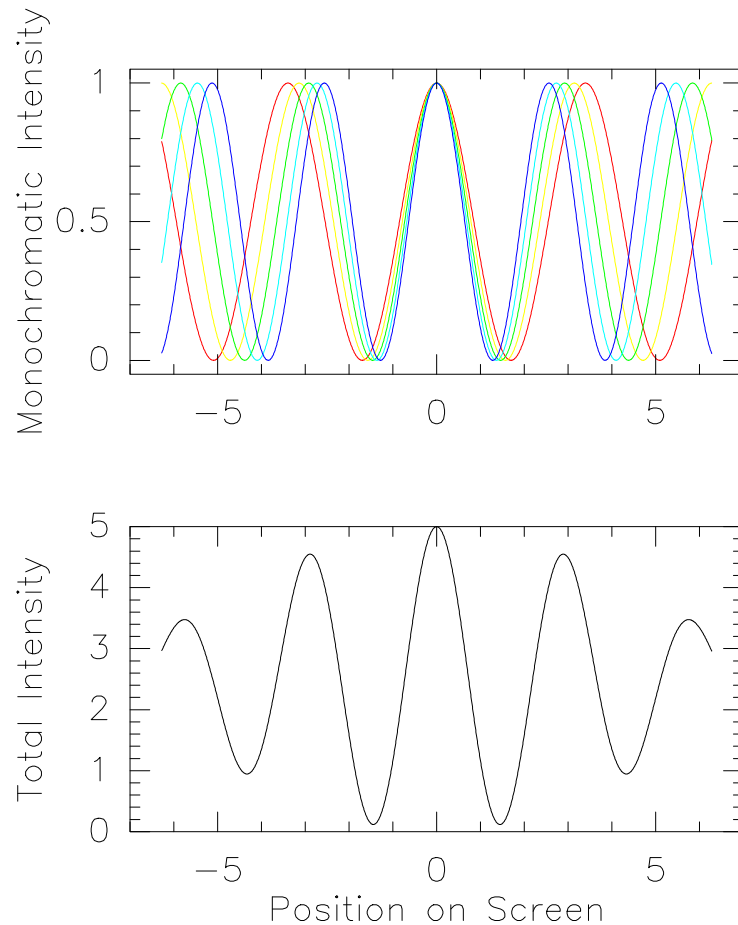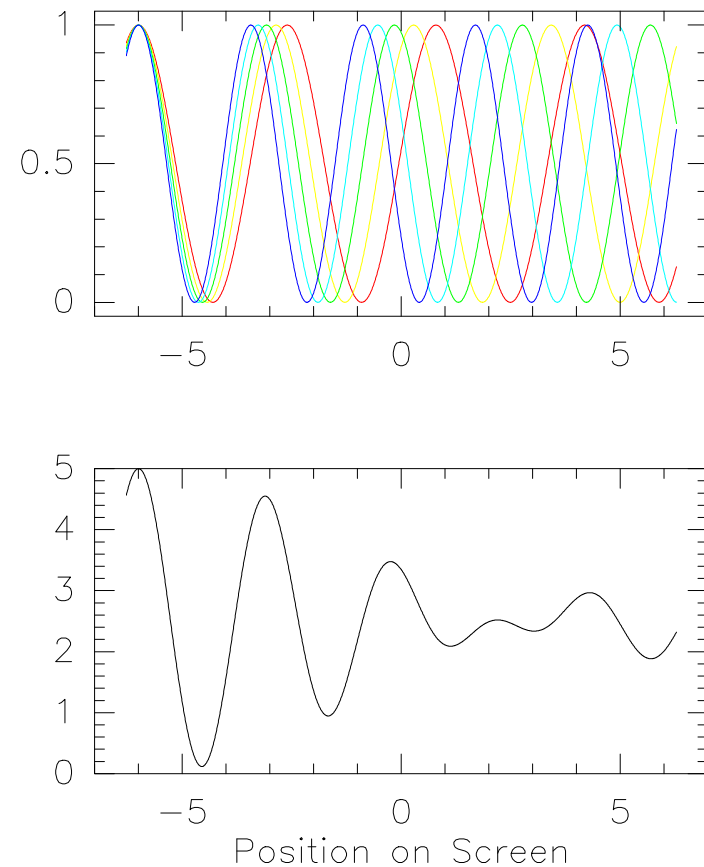⇒ Translation of the fringe pattern.
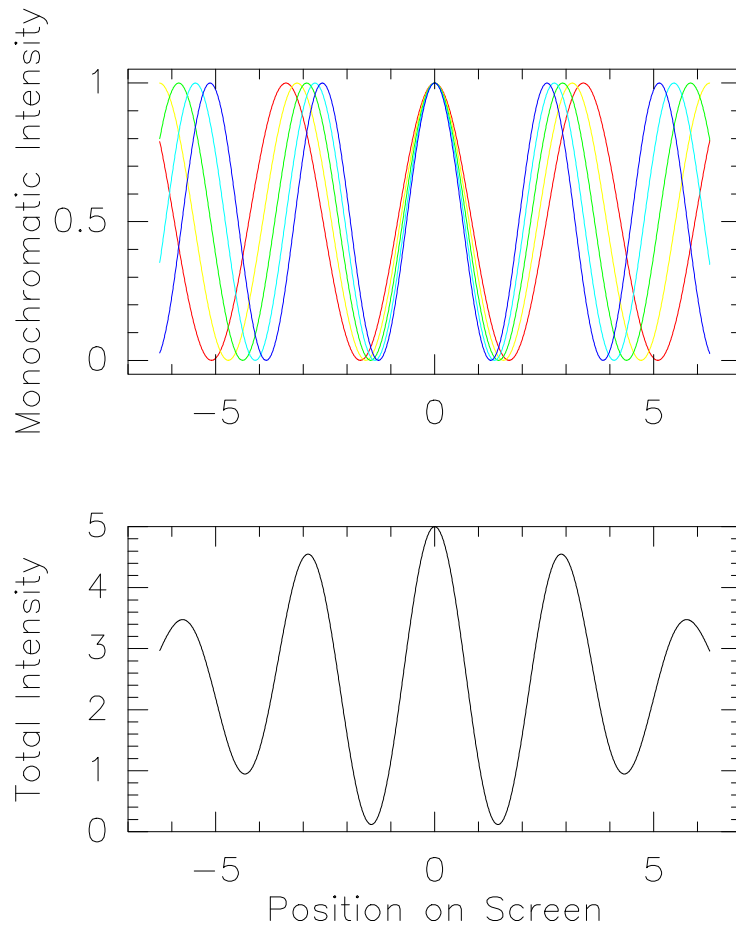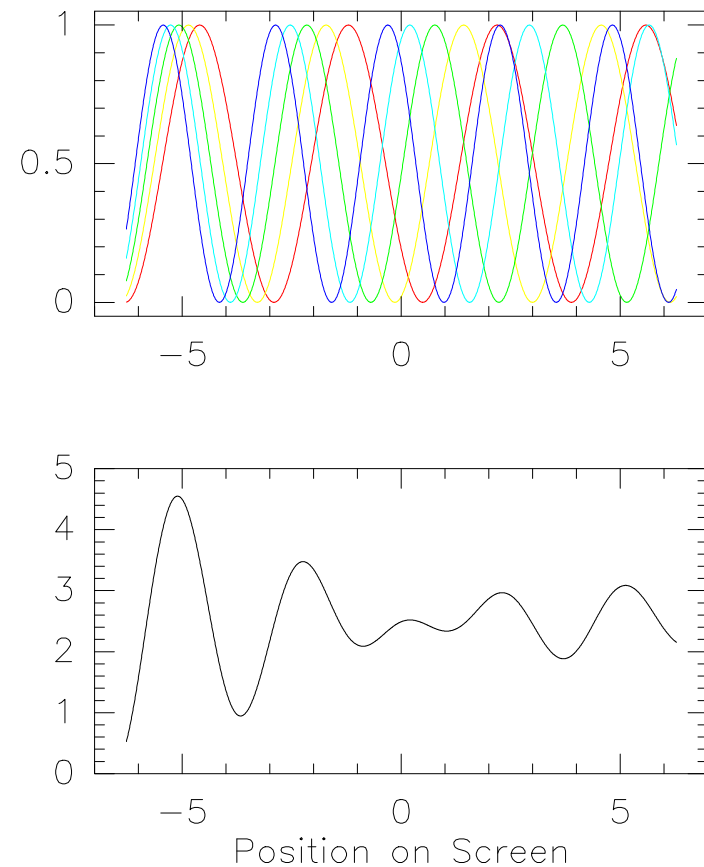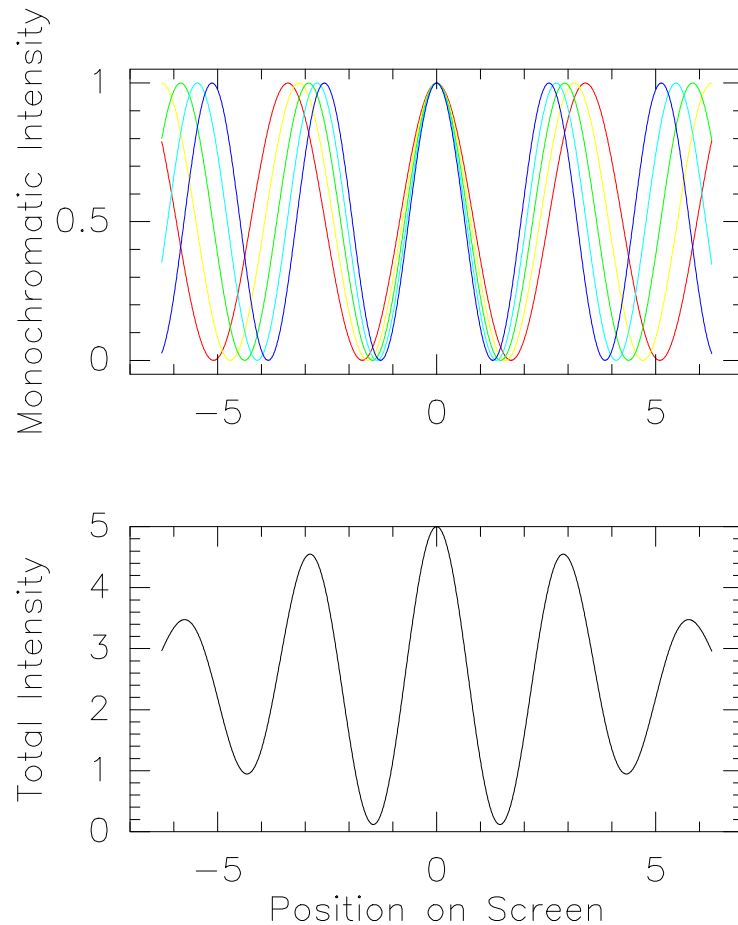⇒ Fringes seem to disappear.

Optic vs Radio Interferometer: I. Measurement Method

		Optic	Radio
Detector	Kind	Quadratic	Linear (Heterodyne)
	Observable	$I = |EE^\star|$	$|E| \exp(i\psi)$
Measure	Method	Optical fringes	Electronic correlation
	Quantity	$|C| = \dfrac{I_{\max} - I_{\min}}{I_{\max} + I_{\min}}$	$|V| \exp(i\phi_V) = <E_1 . E_2>$
Interferometer kind		Additive	Multiplicative

$$I_1 + I_2 + 2\sqrt{I_1 I_2}\,|C| \cos\left(\frac{bx}{\lambda} + \phi_C\right)$$
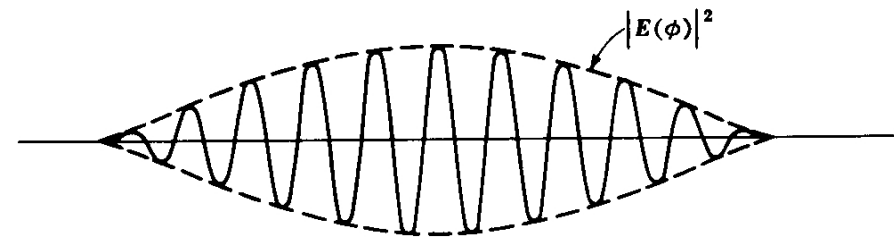
$$\underbrace{|E_1|\,|E_2|\,|C|}_{|V|} \cos\left(\frac{bx}{\lambda} + \overbrace{\phi_C}^{\phi_V}\right)$$

(Heterodyne: lecture by R. Lucas)

Optic vs Radio Interferometer: I. Measurement Method

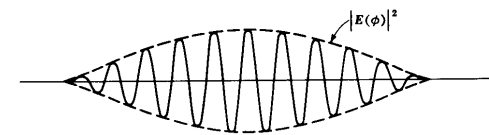|		Optic	Radio
Detector	Kind	Quadratic	Linear (Heterodyne)
	Observable	$I = |EE^\star|$	$|E|\exp(i\psi)$
Measure	Method	Optical fringes	Electronic correlation
	Quantity	$|C| = \frac{I_{\max}-I_{\min}}{I_{\max}+I_{\min}}$	$|V|\exp(i\phi_V) = <E_1.E_2>$
Interferometer kind		Additive	Multiplicative

Multiplicative Interferometer

Avantage: all offsets are irrelevant $\Rightarrow$ Much easier;

Inconvenient: Radio interferometer = bandpass instrument;

$\Rightarrow$ Low spatial frequencies are filtered out.

SD

Multiplicative Interferometer

0m   7.5m   15m   22.5m   30m $\sqrt{u^2 + v^2}$

Optic vs Radio Interferometer: II. Atmospheric Influence
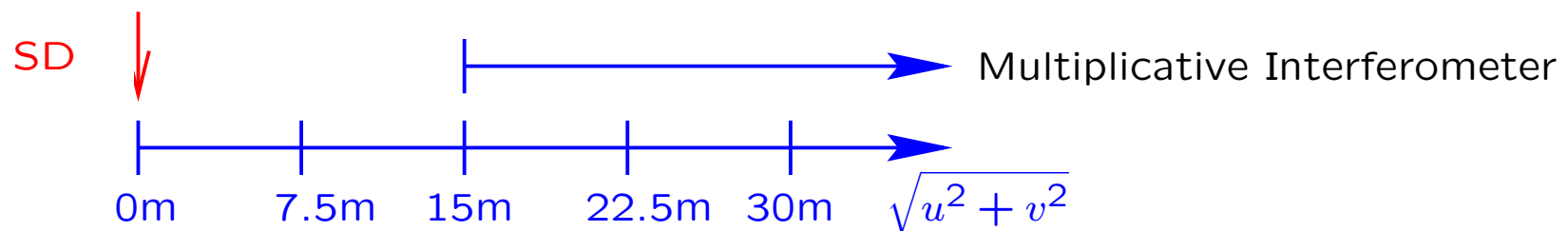
Atmosphere emits and absorbs: (Lecture by J.-M. Winters).
Signal = Transmission * Source + Atmosphere.

- Optic: $\begin{cases} \text{Source} \gg \text{Atmosphere} \\ \text{Transmission} \sim 1 \end{cases} \Rightarrow$ transparent;

- Radio: $\begin{cases} \text{Source} \ll \text{Atmosphere} \\ \text{Transmission can be small} \end{cases} \Rightarrow$ fog.

Good news: Atmospheric noise uncorrelated

$\Rightarrow$ Correlation suppresses it!

Bad news: Transmission depends on weather and frequency.

$\Rightarrow$ Astronomic sources needed to calibrate the flux scale!

(Lectures by R. Lucas and F. Gueth)

Atmosphere is turbulent: $\Rightarrow$ Phase noise (Lecture by A. Karastergiou).
Timescale of atmospheric phase random changes:

- Optic: 10-100 milli secondes;

- Radio: 10 minutes.
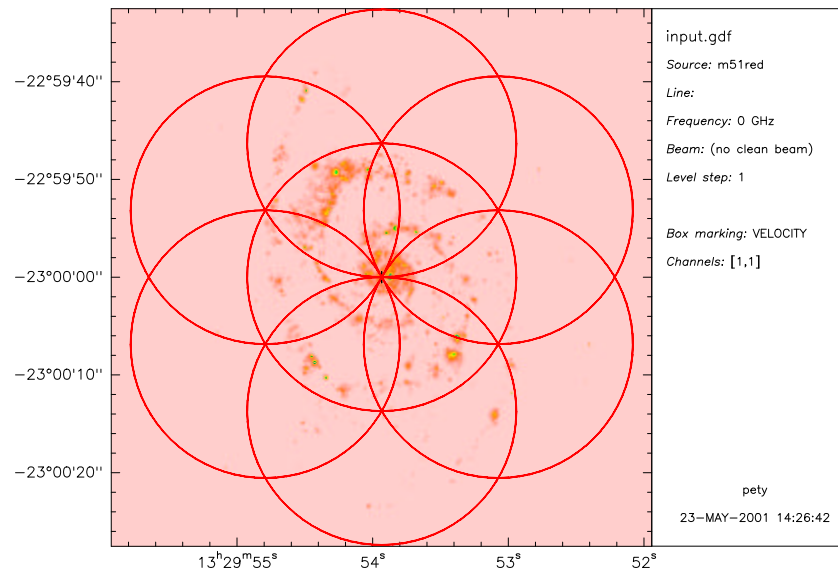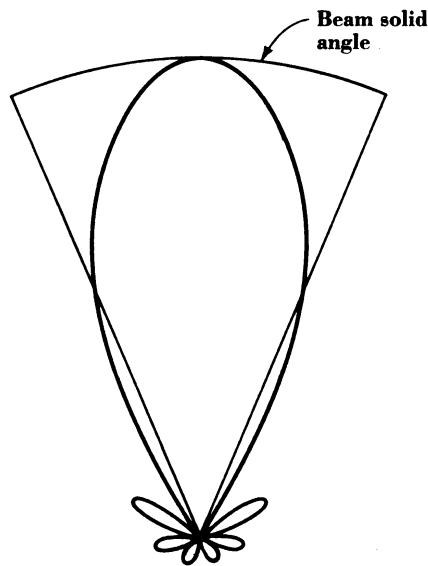
$\Rightarrow$ Radio permits phase calibration on a nearby point source (*e.g.* quasar).

Instantaneous Field of View

One pixel detector:

- Single Dish: one image pixel/telescope pointing;

- Interferometer: numerous image pixels/telescope pointing
  - Field of view = Primary beam size;
  - Image resolution = Synthesized beam size.

Wide-field imaging: ⇒ mosaicing (Lecture by F. Gueth).

Conclusion

mm interferometry:

- A bit more of theory;

- Lot's of experimental details (*e.g.* lecture by R. Lucas).

Why caring about technical details: Some of them must be understood to know whether you can trust your data.

By the end of this week, you should be ready to use PdBI!
(Lecture by R. Neri and examples from users)

Bibliography

- "Synthesis Imaging". Proceedings of the NRAO School. R. Perley, F. Schwab and A. Bridle, Eds.

- "Proceedings from IMISS2", A. Dutrey Ed.

- "Interferometry and Synthesis in Radio Astronomy", R. Thompson, J. Moran and G. W. Swenson, Jr.

Photographic Credits

- M. Born & E. Wolf, "Principles of Optics".

- J. W. Goodman, "Statistical Optics".

- J. D. Kraus, "Radio Astronomy".

Lexicon

- Beam: Antenna diffraction pattern.

- Primary Beam: Instantaneous field of view (Single-Dish Beam).

- Synthesized Beam: Image resolution (Interferometer Beam).

- Configuration: Antenna layout of interferometer.

- Baseline: Distance between two antenna.

- uv-plane: Fourier plane.

- Visibilities: $\sim$ Fourier components of the source.

- Fringe stopping: Temporal variation of delay correction needed to avoid translation of the white fringe.

- Heterodyne: Principle of linear detection.

- Correlator: Where visibilities are measured by correlation of signal coming from pairs of antenna.